

Validation of the DESI DR2 measurements of baryon acoustic oscillations from galaxies and quasars

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The Dark Energy Spectroscopic Instrument (DESI) Data Release 2 (DR2) galaxy and quasar clustering data represents a significant expansion of data from Data Release 1 (DR1), providing improved statistical precision in baryon acoustic oscillation (BAO) constraints across multiple tracers, including bright galaxies, luminous red galaxies, emission line galaxies, and quasars. In this paper, we validate the BAO analysis of DR2. We present the results of robustness tests on the blinded DR2 data and, after unblinding, consistency checks on the unblinded DR2 data. All results are compared with those obtained from a suite of mock catalogs that replicate the selection and clustering properties of the DR2 sample. We confirm the consistency of DR2 BAO measurements with DR1 while achieving a reduction in statistical uncertainties due to the increased survey volume and completeness. The combined BAO precision, including both statistical and systematic errors, improves from $\sim 0.52\%$ in DR1 to 0.30% in DR2—a factor of 1.7 gain. We assess the impact of analysis choices, including different data vectors (correlation function vs power spectrum), modeling approaches and systematics treatments, and an assumption of the Gaussian likelihood, finding that our BAO constraints are stable across these variations and assumptions with a few minor

refinements to the baseline setup of the DR1 BAO analysis. We summarize a series of pre-unblinding tests that confirmed the readiness of our analysis pipeline, the final systematic errors, and the DR2 BAO analysis baseline. The successful completion of these tests led to the unblinding of the DR2 BAO measurements, ultimately leading to the DESI DR2 cosmological analysis, with their implications for the expansion history of the Universe and the nature of dark energy presented in the DESI key paper (companion paper).

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I. INTRODUCTION

Baryon acoustic oscillations (BAO) have emerged as one of the most robust and reliable probes for studying the expansion history of the Universe. These oscillations, imprinted in the large-scale distribution of galaxies and quasars, provide a cosmic standard ruler, enabling precise distance measurements across vast cosmic epochs [1]. Their characteristic physical scale is precisely determined by cosmic microwave background (CMB) measurements, allowing BAO to serve as a powerful tool for mapping the Universe's expansion history. Over the past two decades, BAO measurements have become indispensable in cosmology, providing key constraints on the parameters governing the standard cosmological model, including the nature of dark energy [2–4]. The first detection of the BAO peak by the Sloan Digital Sky Survey (SDSS) [5] and the Two-degree Field Galaxy Redshift Survey [6,7] initiated the progress toward BAO becoming a cornerstone of observational cosmology.

Subsequent advancements have further refined the precision of BAO as a cosmological probe, driven by both theoretical and observational progress. On the theoretical side, improvements in BAO reconstruction techniques [8–10] have significantly enhanced our ability to extract the primordial BAO signal. On the observational side, major spectroscopic surveys—including the WiggleZ Dark Energy Survey [11–13], the Baryon Oscillation Spectroscopic Survey (BOSS) [14–16], and the extended BOSS (eBOSS) [17,18]—have steadily improved BAO constraints, solidifying its role as a foundational probe in cosmology. Beyond galaxy and quasar clustering, alternative approaches have been developed to measure BAO at higher redshifts and in different observational regimes. The Ly α forest absorption in quasar spectra provides a means to probe BAO at $z > 2$, as demonstrated in [19,20]. Meanwhile, photometric surveys have enabled transverse BAO measurements [21], though spectroscopic redshifts provide higher precision for a given number of tracers. These complementary techniques extend BAO constraints across a broader redshift range, enhancing our ability to probe cosmic expansion.

The Dark Energy Spectroscopic Instrument (DESI) represents a significant advancement in large-scale structure (LSS) surveys, aiming to map the 3D distribution of galaxies and quasars over an unprecedented volume [22]. Over its 5-year survey (2021–2026), DESI plans to obtain

spectra of about 40 million galaxies and quasars across 14 000 square degrees, covering a redshift range up to $z \sim 3.5$ [23–25]. DESI released its first cosmology analysis in April 2024 using Data Release 1 (DR1) [26], which included BAO measurements from galaxies and quasars [27], Ly α forest [28] and its cosmological interpretation [29]. This was followed by the release of full-shape measurements [30] and their corresponding cosmological analysis [31,32] in November. The DR1 BAO analysis introduced several advancements, including a catalog-level blinded analysis, a unified framework for all tracers, extensive systematic tests, and improvements in reconstruction and modeling. These measurements provided new constraints on the cosmic expansion history, reinforcing the robustness of BAO as a cosmological probe.

Building on this foundation, DESI Data Release 2 (DR2) [33] provides an even more powerful data set, expanding the survey area and achieving higher completeness. These improvements significantly enhance the statistical precision of BAO measurements. The combined precision of BAO constraints across six redshift bins improves from $\sim 0.52\%$ in DR1 to $\sim 0.24\%$ in DR2, more than doubling the measurement precision. However, this increased precision also heightens sensitivity to systematic effects, requiring even more rigorous validation procedures to ensure robust cosmological constraints.

This paper presents the validation process for the DR2 BAO measurements, detailing the unblinding tests, systematic quantifications, and finalization of the baseline analysis setup. While the general framework follows the DR1 BAO analysis, several refinements were introduced based on the tests conducted in this study. A critical aspect of this validation is the blinding scheme, which prevents confirmation bias. We employed the DR1 blinding pipeline [34–36], ensuring that key results remained blinded until all predefined validation steps were completed. The unblinded DR2 BAO constraints and their cosmological implications are presented in the DESI key paper [37].

The structure of this paper is as follows. In Sec. II, we describe the DESI DR2 data set, detailing the sample selection, clustering estimators, and blinding procedure. Section III introduces the mock catalogs used to validate the analysis. Section IV outlines the modeling framework, including BAO reconstruction, parameter inference, and the construction of covariance matrices. Section V presents key methodological updates in DR2 and the treatment of

TABLE I. Statistics for each DESI tracer type and redshift bin used for BAO measurements. Columns show the number of good redshifts, the redshift range, the effective redshift z_{eff} , the survey area, the adopted value of $P_0(k = 0.14)$ for each tracer, and the corresponding effective volume V_{eff} . The effective redshift z_{eff} represents the weighted mean redshift of a sample, indicating where the measurement is most sensitive. The parameter $P_0(k = 0.14)$ is a fiducial power spectrum value used to optimally weight the clustering signal for each tracer. The effective volume V_{eff} quantifies the statistical power of a given tracer sample in measuring BAO, accounting for the survey volume and the signal-to-noise (S/N) contribution from the tracer number density and clustering amplitude. The survey area differs for each tracer due to priority vetoes (e.g., a QSO target can remove sky area from lower-priority samples) and small variations in imaging vetoes. The effective volume is computed following Eq. (2.2) in [27].

Tracer	No. of redshifts	Redshift range	z_{eff}	Area [deg ²]	$P_0(k = 0.14)$	V_{eff} (Gpc ³)
BGS	1 188 526	$0.1 < z < 0.4$	0.295	12 355	7000	3.8
LRG1	1 052 151	$0.4 < z < 0.6$	0.510	10 031	10 000	4.9
LRG2	1 613 562	$0.6 < z < 0.8$	0.706	10 031	10 000	7.6
LRG3	1 802 770	$0.8 < z < 1.1$	0.922	10 031	10 000	9.8
ELG1	2 737 573	$0.8 < z < 1.1$	0.955	10 352	4000	5.8
ELG2	3 797 271	$1.1 < z < 1.6$	1.321	10 352	4000	8.3
QSO	1 461 588	$0.8 < z < 2.1$	1.484	11 181	6000	2.7

systematic uncertainties. Section VI details our main findings, comparing DR1 and DR2 BAO constraints, assessing robustness across multiple tests, and evaluating systematic effects. Section VII explores post-unblinding tests, including correlated systematics and the impact of fiducial cosmology assumptions. Finally, in Sec. VIII, we summarize our findings and confirm the robustness of the DESI DR2 BAO analysis, establishing its readiness for cosmological interpretation in [37].

II. DATA

A. DESI DR2

The DESI DR2 [33] data set represents the culmination of nearly 3 years of observations using the DESI instrument [38,39], from May 14, 2021 to April 9, 2024. Conducted at the Nicholas U. Mayall Telescope on Kitt Peak National Observatory, Arizona, DESI observes the spectra of 5000 targets [40] simultaneously within a 7 deg² field of view [41], utilizing robotic positioners [42] to align optical fibers [43] with celestial coordinates. These fibers channel light to ten climate-controlled spectrographs, enabling precise redshift measurements critical for cosmological studies.

DESI observations use a dynamic time allocation strategy, which divides observing time into “bright time” and “dark time” programs based on observing conditions [44], with distinct target classes defined for each program. This ensures optimal data quality for different target classes, including galaxies, quasars, and stars. The DR2 data set contains 6671 dark-time and 5172 bright-time tiles, each corresponding to specific sky positions and associated target sets. The sky coverage of these tiles can be seen in the top panel of Fig. 2 in our companion paper [37].

In this work, we use LSS catalogs that were constructed based on the results of the DESI spectroscopic reduction [45] and redshift estimation (Redrock [46,47]) pipelines applied to the DR2 data set in a homogeneous processing

run denoted as “Loa.” The DESI LSS catalog pipeline was detailed in [48], with the specific choices (e.g., mask definitions, completeness weights, treatment for imaging systematics) applied to the version used for DESI DR2 BAO measurements mostly matching those described in [49]. All new choices are described in Sec. II A of the companion paper [37].

The LSS catalogs are split into four distinct tracer types that apply different target selection criteria: the Bright Galaxy Sample (BGS [50]), Luminous Red Galaxies (LRGs [51]), Emission Line Galaxies (ELGs [52]), and Quasars (QSOs [53]). For BAO measurements, the samples are split into the same redshift bins as applied to DR1 [49]. Basic details on the sample size in each redshift bin are provided in Table I. Additionally, the data from the third redshift bin of the LRGs and the first one of the ELGs are combined into a single $0.8 < z < 1.1$ LRG + ELG sample, applying weights that optimally balance the contribution from each target type [54]. The DR2 LSS catalogs used in this work will be released publicly with DR2 version v1.1/BAO. Particular details on the characteristics of each sample and how this informs the tests presented throughout this work are introduced in the next subsection.

B. Sample characteristics and splits

The DESI survey includes four primary tracer samples: BGS, LRGs, ELGs and QSOs. Each of these tracers has distinct characteristics in terms of number density, redshift distribution, and clustering amplitude, which influence their roles in cosmological analyses.

BGS: This sample is subselected from the DESI BGS_BRIGHT selection, which is flux-limited at $r < 19.5$. Due to this flux limit, the nominal BGS_BRIGHT sample exhibits a number density that varies significantly with redshift. To mitigate this redshift evolution and obtain a more uniform sample, we apply an absolute magnitude cut of $M_r < -21.35$. This adjustment

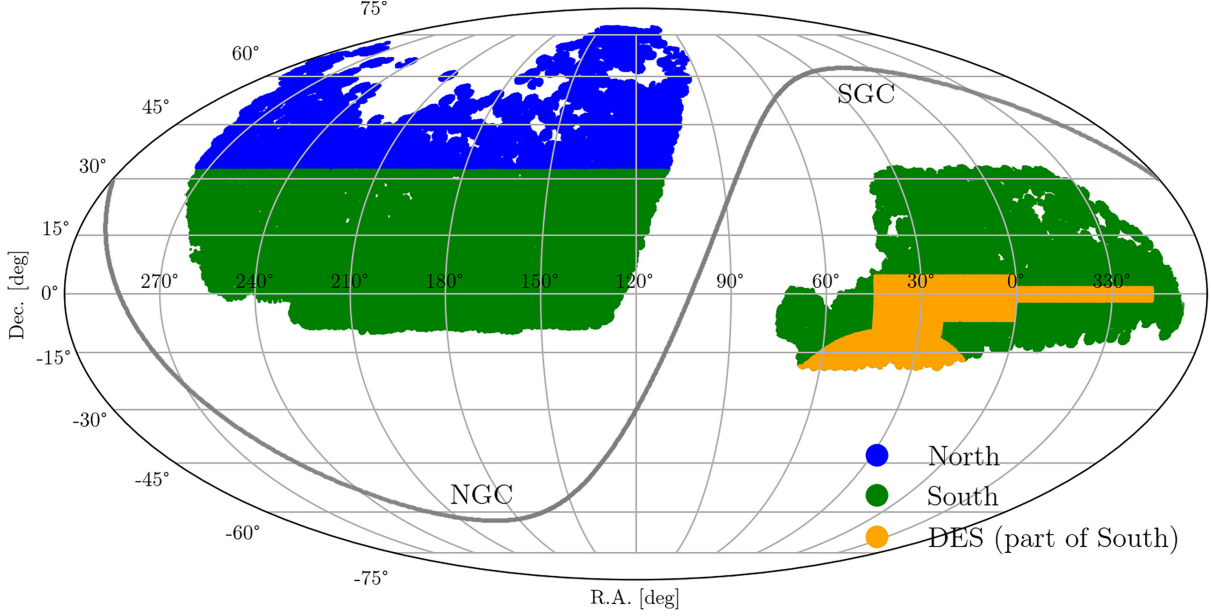


FIG. 1. The DESI DR2 footprint used in the LSS analysis, highlighting different survey regions. The survey is divided into the NGC and SGC, separated by the Galactic equator (gray curve). Within these caps, the imaging footprint is further divided into North (blue; light shading in grayscale) and South (green; medium shading) regions based on the sources of imaging data. The DES region (yellow; dark shading) is a subset of the South region, characterized by deeper imaging and smaller point-spread functions compared with DECaLS (see text). These divisions play a key role in our robustness tests, ensuring that BAO measurements remain consistent across different survey regions.

ensures a more consistent number density across the redshift range. For additional details, see the companion paper [37].

LRG: This sample maintains an approximately constant number density of $5 \times 10^{-4} h^3 \text{ Mpc}^{-3}$ in the redshift range $0.4 < z < 0.8$. Beyond $z = 0.8$, the number density decreases due to a z -band flux limit. Due to its large survey volume and strong clustering signal relative to shot noise, the LRG sample provides the highest effective volume ($V_{\text{eff}} = 9.8 \text{ Gpc}^3$), making it a key contributor to the constraining power of cosmological analyses.

ELG: This sample has a comparable number density to the LRGs for $z = 0.8$,¹ where the LRG number density begins to sharply drop and extends up to $z = 1.6$. However, the clustering amplitude of ELGs is approximately one-third that of LRGs, which reduces their constraining power despite their greater abundance. The ELG sample is also more susceptible to systematic variations in target density caused by imaging systematics. Nevertheless, it has been demonstrated that these systematics do not significantly affect BAO measurements [55]. Imaging systematics are mitigated by applying weights to the galaxy samples to nullify trends with imaging properties.

¹The relative number density between ELGs and LRGs depends on the completeness in a given area, leading to variations in observed values across different regions of the survey footprint.

QSO: The quasar sample is the sparsest among the DESI tracers, with a number density of approximately $2.5 \times 10^{-5} h^3 \text{ Mpc}^{-3}$. This low number density results in the sample being shot-noise-dominated, which limits the precision of clustering measurements. Despite this, QSOs provide valuable information for probing the LSS due to the large cosmic volume spanned by their high redshift range.

1. Spatial and imaging-based splits

Beyond these intrinsic characteristics, the spatial distribution of the samples across the survey footprint introduces additional factors to account for. For all tracer types, there are distinct survey regions of interest, as shown in Fig. 1. These regions are defined by two key divisions:

- (1) The Galactic Cap (GC) Division: separates the North Galactic Cap (NGC) and South Galactic Cap (SGC) based on spatial location relative to the Galactic equator.
- (2) The Imaging-Based Division: defines the North and South regions based on the sources of imaging data used for DESI targeting.

We describe these divisions in more detail below.

The Galactic cap division: The DESI LSS catalogs are divided into NGC and SGC for convenience, as these regions are spatially disjoint and sufficiently separated across the Galactic equator. Due to this separation, clustering measurements can be computed independently for each

region and then combined into a fiducial measurement without loss of information.

The NGC footprint contains 70% of the DR2 bright-time (BGS) sample and 66% of the dark-time (LRGs, ELGs, QSOs) sample. One of our robustness tests evaluates the impact of restricting the clustering analysis to the NGC footprint.

The imaging-based division: The optical imaging data used for DESI targeting comes from two distinct sources, defining the North and South imaging regions:

- (1) The North region contains photometric data from the BASS + MzLS surveys [56]. All NGC data with declination greater than 32.375 degrees is part of the North imaging region (Fig. 1, blue region).
- (2) The South region is entirely observed with DECam [57]. Approximately 1130 deg² in the SGC is targeted by the Dark Energy Survey (DES [58]; Fig. 1, yellow region). The remaining area (≈ 3580 deg² in the SGC and ≈ 5770 deg² in the NGC) is covered by the DECaLS survey [24] (Fig. 1, green region).
- (3) Within the South region, the DES forms a distinct subset. DES typically provides deeper imaging and smaller point-spread functions compared with DECaLS.

The North imaging region covers 31% of the DR2 bright-time footprint and 20% of the dark-time footprint, while the DES region covers 7% of the bright-time footprint and 9% of the dark-time footprint.

Since these imaging differences may introduce systematic effects, our analysis includes robustness tests to evaluate the consistency of BAO measurements across these imaging regions. Thus, we examine the effects of excluding the DES and North regions on the clustering measurements across all samples. See Ref. [48] for more details on how the LSS catalogs are processed accounting for the details of these imaging regions.

2. Intrinsic galaxy property splits

In addition to region-based splits, we perform robustness tests based on two types of intrinsic properties: *stellar mass splits* and *magnitude splits*. These tests help assess potential biases in BAO measurements arising from the dependence of galaxy clustering on stellar mass and luminosity, as both properties are correlated with galaxy bias, which affects how galaxies trace the underlying LSS.

Stellar mass splits are based on estimates from the stellar mass catalog described in Appendix C of [51]. Stellar masses are derived using $g-r$, $r-z$, $z-W_1$, and W_1-W_2 colors from the Legacy Survey DR9, processed with a random forest model [59] trained on the S82-MGC galaxy sample [60]. A key update from [51] is the use of kibo-v1 spectroscopic redshifts instead of photometric redshifts, significantly improving the accuracy of the mass estimates. For LRGs with DESI redshifts, the stellar mass uncertainty calibrated against the training sample is $\sigma_{\text{NMAD}} = 0.086$,

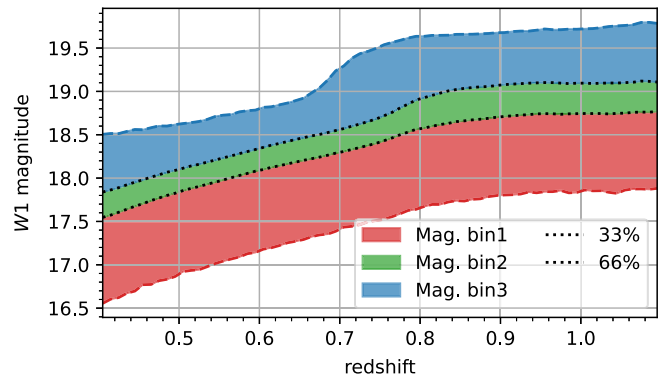


FIG. 2. Illustration of the magnitude-based sample split applied to the DESI LSS galaxy samples. The plot shows the W_1 -band magnitudes as functions of redshift, with dashed lines indicating the 33% and 66% percentile thresholds used to define three magnitude bins. This binning strategy helps assess the impact of galaxy brightness on clustering measurements while ensuring a consistent redshift distribution within each bin.

making them well suited for mass-split analyses. BGS is likely to be similar, so we perform stellar mass splits for BGS and LRGs. In contrast, the scatter for ELGs is significantly larger, reducing the reliability of their stellar mass estimates; thus, we do not apply stellar mass splits to the ELG sample. We divide the BGS sample into two stellar mass bins while dividing the LRG sample into three bins.

Magnitude splits are applied to tracers with sufficiently high number densities, including the BGS, LRG and ELG samples.

For BGS, LRGs and ELGs, magnitude splits are performed using extinction-corrected r -band, W_1 -band and g -band magnitudes, respectively. We divide the BGS and ELG samples into two magnitude bins while dividing the LRGs into three bins. These splits help to probe the clustering dependence on galaxy brightness, particularly in cases like ELGs where stellar mass splits are unreliable.

To ensure a consistent redshift distribution across subsamples in all stellar mass and magnitude splits, we apply percentile-based thresholds within small redshift intervals. Specifically, for the BGS and ELG samples, we use the median (50%) values of the stellar mass and extinction-corrected magnitude in each $\Delta z = 0.01$ interval to define the splits. For LRGs, we refine this approach by using the 33% and 66% percentiles, dividing the sample into three bins, as shown in Fig. 2. Additionally, the same binning strategy is applied to the corresponding random catalogs to maintain a matched spatial and redshift distribution, ensuring that systematic effects do not bias the clustering measurements.

The primary motivation behind these splits is to assess and mitigate any potential biases in the DESI DR2 BAO measurements. Spatial and imaging-based splits help evaluate the impact of survey selection effects, while intrinsic property splits provide insight into galaxy bias variations within the same sample. By performing these

robustness tests, we ensure that our DESI DR2 BAO constraints remain unbiased and reliable for cosmological interpretation.

C. Clustering measurements

To quantify the LSS of the Universe, we measure the two-point correlation function and the power spectrum, which provide complementary descriptions of galaxy clustering. The correlation function characterizes clustering in configuration space (CS), while the power spectrum offers a Fourier-space (FS) perspective, with both methods capturing the same underlying information but differing in sensitivity to systematics and scale-dependent features. To maintain consistency with previous analyses, we follow the same code settings as in DR1 [49], applying well-tested estimators to both correlation function and power spectrum measurements. Additionally, various observational weights are applied to correct for systematics, including imaging systematics weights, redshift failure weights, and the Feldman-Kaiser-Peacock (FKP) weight, which optimally downweights galaxies in high-density regions to reduce sample variance [61]. These weights are consistently incorporated in both configuration-space and FS clustering measurements.

1. Correlation function estimator

We use the Landy-Szalay estimator [62] to compute the two-point correlation function, which measures the excess probability of finding two galaxies separated by a distance s and cosine of the angle relative to the line of sight, μ . The estimator is defined as

$$\hat{\xi}(s, \mu) = \frac{DD(s, \mu) - 2DR(s, \mu) + RR(s, \mu)}{RR(s, \mu)}, \quad (1)$$

where

- (1) $DD(s, \mu)$ represents the weighted number of galaxy-galaxy pairs.
- (2) $DR(s, \mu)$ is the number of galaxy-random pairs.
- (3) $RR(s, \mu)$ denotes the number of random-random pairs.

The random catalog is constructed to match the survey footprint and the selection function to mitigate systematic effects. From the correlation function, we compute multipole moments (monopole $\ell = 0$, quadrupole $\ell = 2$, hexadecapole $\ell = 4$) using Legendre polynomials:

$$\hat{\xi}_\ell(s) = \frac{2\ell + 1}{2} \int_{-1}^1 d\mu \hat{\xi}(s, \mu) \mathcal{L}_\ell(\mu). \quad (2)$$

For post-reconstruction measurements (Sec. IV B), we use a modified version of the Landy-Szalay estimator, as in [63]. We perform the correlation function measurements using a modified version of the CorrFunc pair counting

code [64], implemented in PYCORR.² We use a bin width of $4 h^{-1}$ Mpc in s and 240μ bins from -1 to 1 .

2. Power spectrum estimator

In FS, we estimate the power spectrum using the FKP estimator [61], which accounts for the effects of survey geometry. The weighted galaxy fluctuation field is given by

$$F(\mathbf{r}) = n_d(\mathbf{r}) - \alpha n_r(\mathbf{r}), \quad (3)$$

where $n_d(\mathbf{r})$ and $n_r(\mathbf{r})$ are the weighted galaxy and random number densities, with the latter having a total weighted number $1/\alpha$ times that of the data catalog. This expression is appropriately modified for post-reconstruction measurements as in [65].

Power spectrum multipoles are computed using the Yamamoto estimator [66], which efficiently handles line-of-sight variations across the survey volume:

$$\hat{P}_\ell(k) = \frac{2\ell + 1}{AN_k} \sum_{\vec{k} \in k} \sum_{\vec{r}_1} \sum_{\vec{r}_2} F(\vec{r}_1) F(\vec{r}_2) \mathcal{L}_\ell(\hat{k} \cdot \hat{\eta}) \times e^{i\vec{k} \cdot (\vec{r}_2 - \vec{r}_1)} - \mathcal{N}_\ell. \quad (4)$$

The summations extend over all galaxy pairs with positions $\vec{r}_1$ and $\vec{r}_2$, as well as over wave vectors $\vec{k}$ within the given k bin. $\hat{\eta}$ is the line-of-sight direction. $\mathcal{N}_\ell$ accounts for the shot-noise correction applied to the monopole component, while A serves as the normalization factor. N_k corresponds to the total number of modes contributing to the estimation in the given k bin.

To efficiently implement the Yamamoto estimator using fast Fourier transforms, we follow the approach described in [67], which allows for rapid computation of the power spectrum multipoles in large-volume surveys. This methodology is incorporated into the NBODYKIT framework [68] and is further optimized in PYPOWER,³ which we use for our analysis.

We follow the same code settings as in DR1 [49]. While our fiducial results are based on the correlation function, following the choice of DR1 based on the performance of the analytical covariance matrix, the power spectrum analysis serves as a cross-check, and its results are presented in Appendix A.

D. Blinding

Blinding is an integral part of the validation process, designed to prevent confirmation bias. Following the DESI DR1 blinding scheme [34], DESI DR2 BAO measurements were kept blinded during the validation process, including the determination of the systematic error budget (see

²<https://github.com/cosmodesi/pycorr/>.

³<https://github.com/cosmodesi/pypower>.

Sec. VB). Unblinding was only performed after completing all predefined validation tests listed in Table VI. For completeness, we briefly summarized the blinding scheme; for further details see [34–36].

We adopt a catalog-level blinding that modifies galaxy redshifts and weights to mimic the effects of a different underlying cosmology. This approach allows collectively blinding three key observables: BAO, redshift-space distortions (RSDs) and primordial non-Gaussianity (PNG).

- (1) *Blinding the expansion rate*: Galaxy redshifts are modified by first converting the observed redshifts to comoving distances using a blind cosmology, then reconverting them to blinded redshifts based on the fiducial cosmology. This process introduces an unknown dilation of all scales, in both the radial and transverse directions. As a result, both the isotropic and anisotropic BAO signals are blinded, ensuring that the inferred expansion rate remains concealed during validation.
- (2) *Blinding RSD (RSD shift)*: Galaxy redshifts are perturbed based on the local density and peculiar velocity field, mimicking changes to the growth rate of structure. This distorts anisotropies in the power spectrum, therefore blinding RSD.
- (3) *Blinding PNG (scale-dependent bias)*: The scale-dependent bias signature of PNG [69] is mimicked by applying an additional weight to each galaxy. This weight is computed from the real-space density field (δ^r), reconstructed from the observed galaxy distribution, and incorporates a blinded value of f_{NL} . This modification induces a scale-dependent bias effect, ensuring that PNG constraints remain blinded in the large-scale power spectrum.

We adopt the *Planck* 2018 results [70] as our fiducial cosmology. The blind cosmology is randomly selected following the methodology outlined in [34], introducing small shifts in the BAO scaling parameters, the growth of structure, and the PNG parameter f_{NL} . These controlled modifications ensure that cosmological constraints remain blinded throughout validation, preventing any confirmation bias in the analysis. The fiducial parameters are

$$\begin{aligned} \omega_b &= 0.02237, & \omega_{\text{cdm}} &= 0.12, & h &= 0.6736, \\ A_s &= 2.083 \times 10^{-9}, & n_s &= 0.9649, & N_{\text{ur}} &= 2.0328, \\ N_{\text{ncdm}} &= 1.0, & \omega_{\text{ncdm}} &= 0.0006442, & w_0 &= -1, & w_a &= 0. \end{aligned}$$

Here, $\omega_b = \Omega_b h^2$ and $\omega_{\text{cdm}} = \Omega_{\text{cdm}} h^2$ are the physical baryon and cold dark matter densities, where Ω_b and Ω_{cdm} represent the corresponding density parameters. A_s and n_s are the amplitude and tilt of the primordial power spectrum, N_{ur} is the effective number of relativistic species, N_{ncdm} and ω_{ncdm} are the number and physical density of massive neutrinos, and w_0 and w_a give the present-day value and

time evolution of the dark energy equation of state under the Chevallier-Polarski-Linder parametrization [71,72]. These values define the baseline cosmological model used in our analysis.

The final covariance matrices for the BAO analysis were computed after unblinding by applying the same predefined procedure to the unblinded catalogs and clustering measurements. This step was necessary because accurately calibrating the covariance matrices requires knowledge of the actual clustering signal in the data. While the blinding procedure ensured that all systematic and methodological choices were made independently of the final BAO results, refining the covariance after unblinding allowed us to incorporate the actual statistical properties of the DESI DR2 data set. This ensured that the final measurements remained robust and appropriately reflected the uncertainties in the analysis.

III. MOCKS

Mock catalogs are synthetic data sets that simulate the large-scale distribution of galaxies and quasars in the universe based on theoretical models. These mocks are designed to closely replicate real observations, reproducing the expected clustering statistics, survey geometry, and observational effects. By incorporating known cosmological parameters and realistic noise, mocks provide a controlled environment to test analysis pipelines, evaluate systematic uncertainties, and ensure the robustness of measurements. They play a critical role in LSS surveys like DESI, where the precision of the data demands rigorous validation before drawing scientific conclusions.

For the DESI DR2 BAO analysis, we relied on the AbacusSummit second-generation mocks (Abacus-2 DR2 mocks). The “cutsy” mocks we used, which include angular sky coordinates and redshifts over the entire (planned) DESI footprint, are the same as those used for DR1 analyses and were described in [49]. For this analysis, we have applied the “altmtl” method described in [49,73] to match them to the DESI DR2 footprint and observational completeness. We then passed the outputs through the DESI LSS pipeline to match the selection properties of the different target samples, following the kibo-v1 specification. Percent-level adjustments to the assumed redshift failure fractions were applied to better match the final $n(z)$ to the DR2 data. A more in-depth description of the methodology will be provided in [74].

The input boxes for these mocks are based on halo catalogs from the AbacusSummit suite of base simulations [75], populated with galaxies using a halo occupation distribution (HOD) framework, following the methodology of [76] for dark-time tracers and [77] for BGS. The HOD models used in these mocks were calibrated to the DESI SV3 data, as detailed in [78–80].

For the dark-time tracers, the DR2 mocks are identical to those used in DR1, except for updates in the footprint mask to reflect the DR2 survey geometry.

A detailed comparison between the clustering of these mocks and that of DESI tracers was presented in [49], where it was found that the agreement is at or better than 2% in the inferred real-space overdensity field for all tracers and redshift bins.

The BGS mocks, in contrast, required changes for DR2 related to the fact that we have changed the absolute magnitude selection for the DR2 BGS sample. The BGS mocks include absolute magnitudes and simulate the entire BGS_BRIGHT and BGS_FAINT samples. However, it was found that no absolute magnitude cut applied to the mock BGS_BRIGHT sample could reproduce the observed number density above $z = 0.35$ (whereas the absolute magnitude cut applied to DR1 does produce a good match between data and mock so this was not an issue in DR1). To match the number density (and/or the total number of included redshifts), it was thus required to include data from the mock BGS_ANY sample with a redshift-dependent cut on the absolute magnitude. The resulting sample matches the DR2 sample well in terms of the redshift distribution (and total number), but has a $\sim 10\%$ higher clustering amplitude than observed in the data. While this discrepancy does not affect the BAO scale measurement—since the amplitude difference is absorbed into the free bias parameters of the BAO model—it is an important consideration for studies beyond BAO. A detailed examination of the clustering comparison between BGS mocks and DR2 data is ongoing, and upcoming analyses will incorporate improved mocks that better match the observed data, addressing previous discrepancies. A full description of these updated mocks will be presented in a forthcoming study.

The mock catalogs enabled comparisons between the observed clustering and theoretical expectations, allowing us to validate the reconstruction methodology and assess systematic uncertainties in the BAO fitting process. A total of 25 realizations were analyzed for all tracers. These realizations were created from 25 realizations of simulation boxes with side lengths of $2 h^{-1}$ Gpc. For the BGS mocks, the entire DR2 sample fits within the simulation volume. However, for the dark-time tracers, replication of the box is required. The replicated boxes are the same as those used in DR1, which simply tiled the boxes $3 \times 3 \times 3$ to produce boxes with side lengths $6 h^{-1}$ Gpc. Thus, some added correlation is expected both for the clustering measurements within any redshift bin and between redshift bins. This may decrease the expected precision of the results from any individual realization (as the unique volume should determine this) and does decrease our ability to consider the results from all redshift bins together in an ensemble sense. The results from the mocks remain highly valuable, however, for testing for any biases in the recovery of BAO parameters and examining where DESI results fall within the distribution of the 25 mocks. Even with the

replication issues, recovering a result that is within the distribution implies it is at least a few percent likely.

The mocks provided an essential benchmark for exploring potential systematic effects, ensuring that the final results were unbiased.

IV. MODELING

Accurately extracting the BAO signal from the clustering of galaxies and quasars requires a comprehensive modeling framework that accounts for various physical and observational effects. This process involves combining theoretical predictions, observational data, and statistical methods to interpret the observed signal while minimizing systematic biases. Key components of the modeling include the treatment of geometric distortions, the impact of the comoving sound horizon as a standard ruler, and the role of reconstruction in mitigating nonlinear effects.

To capture the dependence of the observed clustering on these effects, the signal is modeled in terms of dilation parameters that encapsulate the effects of geometric distortions and sound horizon rescaling. These parameters, defined in the subsequent sections, allow the BAO feature to be extracted and interpreted in a manner that is robust to the choice of fiducial cosmology.

A. BAO dilation parameters

In spectroscopic surveys like DESI, angular positions and redshifts are used to infer the comoving coordinates of galaxies and quasars, assuming a fiducial cosmology. If the true cosmology differs from the fiducial one, the reconstructed coordinates will exhibit both isotropic and anisotropic distortions, the latter quantified by the Alcock-Paczynski (AP) effect [81]. These distortions cause the observed modes along and perpendicular to the line of sight to be remapped as

$$k_{\parallel}^{\text{true}} = \frac{D_{\text{H}}^{\text{fid}}(z)}{D_{\text{H}}(z)} k_{\parallel}, \quad \mathbf{k}_{\perp}^{\text{true}} = \frac{D_{\text{M}}^{\text{fid}}(z)}{D_{\text{M}}(z)} \mathbf{k}_{\perp}, \quad (5)$$

where $D_{\text{H}}(z) = c/H(z)$ is the Hubble distance, and $D_{\text{M}}(z)$ is the comoving angular diameter distance, both evaluated at a redshift z .

The template power spectrum is the other key aspect to consider. In the BAO fitting framework, rather than constructing a power spectrum model for each different cosmology we test, we assume that the BAO pattern can be mapped by a simple rescaling of a given fiducial template power spectrum to match the BAO feature in the observed power spectrum: $k' = \frac{r_{\text{d}}}{r_{\text{d}}^{\text{temp}}} k^{\text{true}}$.⁴

⁴Here, we use the notation $r_{\text{d}}^{\text{temp}}$ rather than $r_{\text{d}}^{\text{fid}}$ to distinguish the fiducial sound horizon assumed in the template cosmology from that used in the grid cosmology, which maps sky coordinates to physical distances. For simplicity, we now refer to both collectively as the fiducial cosmology, denoted as “fid,” following the nomenclature used in the DESI key paper [37].

Connecting both points motivates the introduction of two free dilation parameters, which rescale the BAO portion of the power spectrum along and across the line of sight, as $\mathbf{k} \rightarrow (\mathbf{k}_\perp/\alpha_\perp, k_\parallel/\alpha_\parallel)$, and leads to their physical interpretation

$$\alpha_\parallel = \frac{D_H(z)/r_d}{D_H^{\text{fid}}(z)/r_d^{\text{temp}}}, \quad \alpha_\perp = \frac{D_M(z)/r_d}{D_M^{\text{fid}}(z)/r_d^{\text{temp}}}. \quad (6)$$

Alternatively, as is done in this work, these parameters can be translated into an isotropic/anisotropic basis

$$\alpha_{\text{iso}} = (\alpha_\parallel \alpha_\perp^2)^{1/3}, \quad \alpha_{\text{AP}} = \frac{\alpha_\parallel}{\alpha_\perp}, \quad (7)$$

leading to a coordinate transformation of the form

$$k' = \frac{\alpha_{\text{AP}}^{1/3}}{\alpha_{\text{iso}}} \left[1 + \mu^2 \left(\frac{1}{\alpha_{\text{AP}}^2} - 1 \right) \right]^{1/2} k, \quad (8)$$

$$\mu' = \frac{\mu}{\alpha_{\text{AP}} \left[1 + \mu^2 \left(\frac{1}{\alpha_{\text{AP}}^2} - 1 \right) \right]^{1/2}}. \quad (9)$$

Here the unprimed coordinates represent the *observed* coordinates and the primed coordinates (k' , μ') are the coordinates at which the model is evaluated. With this, the BAO portion of the observed galaxy power spectrum, denoted by $P_{\text{w,obs}}(k, \mu)$, is modeled using a rescaled template evaluated at the transformed coordinates (k' , μ'):

$$P_{\text{w,obs}}(k, \mu) = P_{\text{w,mod}} \equiv \frac{1}{\alpha_{\text{iso}}^3} P_{\text{w}}(k', \mu'; \alpha_{\text{iso}}, \alpha_{\text{AP}}). \quad (10)$$

Here, $P_{\text{w,obs}}$ refers to the measured power spectrum, while $P_{\text{w,mod}}$ corresponds to the theoretical template-based prediction after applying the dilation parameters.

Note that α_{iso} and α_{AP} therefore relate to the distances as

$$\alpha_{\text{iso}} = \frac{D_V(z)/r_d}{D_V^{\text{fid}}(z)/r_d^{\text{temp}}}, \quad \alpha_{\text{AP}} = \frac{D_H(z)/D_H^{\text{fid}}(z)}{D_M(z)/D_M^{\text{fid}}(z)}, \quad (11)$$

where $D_V(z)$ is the spherically averaged distance defined as

$$D_V(z) = [z D_H(z) D_M^2(z)]^{1/3}. \quad (12)$$

B. Reconstruction

Reconstruction techniques aim to recover the linear BAO signal from nonlinear degradations caused by gravitational growth, peculiar velocities (RSD), and galaxy bias [82]. This process involves displacing galaxies along the inferred large-scale displacement field, effectively “undoing” the effects of bulk flows. By mitigating these nonlinear degradations, the reconstructed density field not only

enhances the amplitude of the BAO peak, bringing it closer to the linear case and improving measurement precision but also helps remove nonlinear shifts in the location of the BAO feature that could potentially bias the cosmological constraints [82,83].

Under the plane-parallel approximation and assuming the linear continuity equation, the displacement field can be derived from the galaxy overdensity field $\delta_g(\mathbf{k})$ as [84]

$$\mathbf{D}(\mathbf{k}) = -\frac{i\mathbf{k}}{k^2} \frac{S(k)\delta_g(\mathbf{k})}{b(1 + \beta\mu^2)}, \quad (13)$$

where $S(k)$ is a smoothing kernel applied to filter out small scales dominated by highly nonlinear dynamics and shot noise, b is the linear galaxy bias, f is the growth rate of structure, and $\beta = f/b$. The term $b(1 + \beta\mu^2)$ accounts for the contributions of linear bias and large-scale RSDs in the observed galaxy density field and ensures that the derived displacement field reflects the underlying matter displacement field in real space.

In practice, the plane-parallel approximation is not valid when analyzing survey data, where we cannot assume the same line of sight for all pairs of objects. However, Eq. (13) serves to illustrate the point that the displacement calculation requires knowledge of the galaxy bias and the linear growth rate of structure, as well as prescriptions for defining the smoothing kernel and for estimating the galaxy overdensity field from the discrete galaxy positions.

Beyond the plane-parallel approximation, different reconstruction algorithms offer numerical solutions that allow for the estimation of the displacement field for varying lines of sight. As in DR1, our default choice of algorithm for the DR2 clustering analysis is the iterative FFT reconstruction [85], which solves the redshift-space linearized continuity equation in FS by iteratively removing RSDs. This algorithm has been shown to be robust and efficient when compared against other popular algorithms in the literature [84].

We reconstruct the DR2 catalogs using PYRECON,⁵ adopting the same baseline settings as in DR1, which were stress tested on DESI mock catalogs in [86]. The overdensity field is painted on a grid with a cell size of $4 h^{-1}$ Mpc, and is then smoothed by a Gaussian kernel with characteristic width Σ_{sm} , which is taken to be $15 h^{-1}$ Mpc (BGs, LRGs, ELGs) or $30 h^{-1}$ Mpc (QSOs), as summarized in Table II. The linear galaxy bias was determined by [87,88], and corresponds to 1.5 for BGs, 2.0 for LRGs, 1.2 for ELGs, and 2.1 for QSOs. The growth rate of structure is predicted from our fiducial cosmology at the effective redshift of each tracer. We also studied the impact of varying both the bias and the growth rate of structure on the reconstruction process. Reference [89] studied the impact of the choice of fiducial cosmology on the reconstruction

⁵<https://github.com/cosmodesi/pyrecon>.

TABLE II. Reconstruction parameters for each tracer, including the smoothing scale Σ_{sm} , linear galaxy bias b , and growth rate of structure f at the effective redshift z_{eff} .

Tracer	$\Sigma_{\text{sm}} [h^{-1} \text{ Mpc}]$	b	$f(z_{\text{eff}})$
BGS	15	1.5	0.69
LRG	15	2.0	0.83
LRG + ELG	15	1.6	0.86
ELG	15	1.2	0.90
QSO	30	2.1	0.93

process, and the propagation of this effect on the BAO constraints is included in the systematic error budget.

We performed a series of robustness tests for reconstruction, where we added a redshift padding to the galaxy catalogs before reconstructing them (allowing for a more appropriate estimation of the velocity field at the survey edges), and using a redshift-dependent value for the growth rate of structure and the linear galaxy bias. In all cases, we found minimal effects on the resulting clustering measurements and BAO fits.

C. BAO model

With all the necessary components in place, we can now construct the BAO model used to fit the observations. The galaxy power spectrum in redshift space can be modeled as a combination of smooth and oscillatory components to capture the BAO signal. Following [83], the model for the galaxy power spectrum is expressed as

$$P_g(k, \mu) = B(k, \mu)P_{\text{nw}}(k) + C(k, \mu)P_w(k) + D(k), \quad (14)$$

where

- (1) $P_{\text{nw}}(k)$ and $P_w(k)$ are the smooth (no-wiggle) and oscillatory (wiggle) components of the linear power spectrum $P_{\text{lin}}(k)$, respectively, obtained using the *peak average* method [90],
- (2) $B(k, \mu)$ encapsulates the broadband clustering shape, incorporating the effects of RSD and galaxy bias,
- (3) $C(k, \mu)$ isolates the BAO feature while accounting for anisotropic damping caused by nonlinear growth and peculiar velocities, and therefore $C(k, \mu)P_w(k)$ is the only term that is diluted by the $\{\alpha_{\text{iso}}, \alpha_{\text{AP}}\}$, and
- (4) $D(k)$ models any deviations from the linear theory in the broadband shape of the power spectrum multipoles.

The broadband term is expressed as

$$B(k, \mu) = (b_1 + f\mu^2[1 - s(k)])^2 F_{\text{FoG}}, \quad (15)$$

where

- (1) b_1 is the linear galaxy bias,
- (2) f is the linear growth rate of structure,

TABLE III. Gaussian priors on the nonlinear BAO damping parameters used for pre- and post-reconstruction BAO fits across tracers. The values in the table represent the mean (μ) of the Gaussian priors, with uncertainties set to $\pm 1 h^{-1} \text{ Mpc}$ for $\Sigma_{\perp}$ and $\pm 2 h^{-1} \text{ Mpc}$ for $\Sigma_{\parallel}$ and Σ_s , consistent with the priors used in the DR1 BAO analysis [27]. For the combined LRG3 + ELG1 sample, we adopt the same fiducial damping parameters as those used for the LRG sample.

Parameter	Reconstruction	BGS	LRGs	ELGs	QSOs
$\Sigma_{\perp}^{\text{fid}} [h^{-1} \text{ Mpc}]$	Pre	6.5	4.5	4.5	3.5
$\Sigma_{\parallel}^{\text{fid}} [h^{-1} \text{ Mpc}]$	Pre	10.0	9.0	8.5	9.0
$\Sigma_s^{\text{fid}} [h^{-1} \text{ Mpc}]$	Pre	2.0	2.0	2.0	2.0
$\Sigma_{\perp}^{\text{fid}} [h^{-1} \text{ Mpc}]$	Post	3.0	3.0	3.0	3.0
$\Sigma_{\parallel}^{\text{fid}} [h^{-1} \text{ Mpc}]$	Post	8.0	6.0	6.0	6.0
$\Sigma_s^{\text{fid}} [h^{-1} \text{ Mpc}]$	Post	2.0	2.0	2.0	2.0

- (3) $s(k)$ is the smoothing kernel, which is set to zero for unreconstructed catalogs and the `RecSym` reconstruction convention (our default; see [86]), and as $s(k) = \exp[-(k\Sigma_{\text{sm}})^2/2]$ for the `RecIso` convention, where Σ_{sm} represents the smoothing scale used in reconstruction, and
- (4) $F_{\text{FoG}} = [1 + \frac{1}{2}k^2\mu^2\Sigma_s^2]^{-2}$ accounts for the damping from the fingers-of-God (FoG) effect, with Σ_s as the free parameter controlling virial motions.

The oscillatory component is damped anisotropically:

$$C(k, \mu) = (b_1 + f\mu^2)^2 \exp \left[-\frac{1}{2}k^2(\mu^2\Sigma_{\parallel}^2 + (1 - \mu^2)\Sigma_{\perp}^2) \right], \quad (16)$$

where $\Sigma_{\parallel}$ and $\Sigma_{\perp}$ are the damping scales parallel and perpendicular to the line of sight, respectively. We place Gaussian priors on these parameters, with means and uncertainties listed in Table III. These values follow the convention adopted in the DESI DR1 BAO analysis [27], where they were derived in [83] from a combination of theoretical calculations, measurements of the cross-correlation between initial and post-reconstruction density fields in Abacus-2 mocks, and fits to data vectors averaged over many realizations.

The exponential factor models the nonlinear damping of the BAO feature, which reconstruction reduces by mitigating bulk motions. (See Table III for the values used in the pre- and post-reconstruction BAO fits.)

To account for additional broadband contributions beyond the linear theory prediction, we use a spline-based approach to introduce a flexible modeling term:

$$D_{\ell}(k) = \sum_{n=-1}^{n_{\text{max}}} a_{\ell,n} W_3 \left(\frac{k}{\Delta} - n \right), \quad (17)$$

where

- (1) W_3 is a piecewise cubic spline kernel that ensures smooth modeling of broadband deviations,
- (2) Δ is the spacing between spline nodes, chosen to avoid overlap with the BAO signal, typically set to twice the BAO wavelength ($\sim 0.06 h \text{ Mpc}^{-1}$), and
- (3) n_{max} determines the number of spline terms, sufficient to span the analyzed k range ($0.02 h \text{ Mpc}^{-1} < k < 0.30 h \text{ Mpc}^{-1}$) without reproducing the oscillatory features. Based on these criteria, n_{max} is set to 7 in this work, the same as in DR1.

The model multipoles are convolved with the survey window function to ensure accurate comparison with the observed power spectrum. This window function accounts for the effects of the survey geometry and selection function, and is derived following the procedure outlined in [91]. Specifically, since we fit the power spectrum monopole and quadrupole $[P_0(k) \text{ and } P_2(k)]$, rather than the full anisotropic power spectrum $P_{\text{obs}}(k, \mu)$, we integrate over the Legendre polynomials, $\mathcal{L}_\ell(\mu)$, to obtain the model multipoles:

$$P_\ell(k) = \frac{2\ell + 1}{2} \int_{-1}^1 d\mu P_{\text{obs}}(k, \mu) \mathcal{L}_\ell(\mu). \quad (18)$$

To predict the multipoles of the galaxy correlation function in CS, we start from the multipoles of the power spectrum defined in Eq. (14), *excluding* the broadband component $D_\ell(k)$, and perform a Hankel transform:

$$\xi_\ell(s) = \frac{i^\ell}{2\pi^2} \int_0^\infty dk k^2 j_\ell(ks) P_\ell(k), \quad (19)$$

where j_ℓ are the spherical Bessel functions. To account for the remaining broadband contributions, one could in principle Hankel transform the same power spectrum spline functions introduced in Eq. (17). However, as shown in [83], for the fitting range adopted in CS ($60 h^{-1} \text{ Mpc} < s < 150 h^{-1} \text{ Mpc}$), the spline basis functions decay rapidly and contribute negligibly on BAO-relevant scales. We therefore exclude these terms, retaining only the $n = 0, 1$ components of the quadrupole as done in [27].

In addition to these two terms, in the CS BAO fitting we also introduce two additional nuisance parameters for each multipole, aimed to control potential large-scale systematics in the data⁶:

$$\tilde{D}_\ell(s) = b_{\ell,0} + b_{\ell,2} \left(\frac{sk_{\text{min}}}{2\pi} \right)^2, \quad (20)$$

with $k_{\text{min}} = 0.02 h \text{ Mpc}^{-1}$.

⁶In FS, these large-scale systematics can be confined below a certain k_{min} and are therefore more easily mitigated by simply truncating the data vector.

This choice is also physically motivated by the fact that, in FS, one typically removes low- k modes (i.e., $k < k_{\text{min}}$) to avoid poorly constrained or contaminated large-scale modes. The inverse transform of these excluded low- k contributions leads to smooth, long-wavelength distortions in CS, which are naturally captured by an even-power expansion in s . Truncating this expansion at quadratic order yields a broadband model that is flexible enough to absorb large-scale trends, while remaining orthogonal to the oscillatory BAO signal.

D. Covariance matrices

For the DESI DR2 analysis, we rely exclusively on analytical and semianalytical methods to estimate covariance matrices. In doing so, we have saved considerable time and effort that would be needed to create a high-precision mock-based covariance matrix via the calibration, generation and processing of a necessarily large suite of approximate simulations. The faster analytical methods also gave us more flexibility to update the covariance matrices several times as the data evolved, in particular, to keep them blinded initially and to unblind the covariances shortly after the correlation function measurements. The covariance matrix pipelines had been fixed before unblinding; only the input data was changed from blinded to unblinded.

Configuration-space covariances are generated using the RascalC semianalytical code⁷ [92–96], which has already been the fiducial method for the DESI DR1 BAO analysis [27]. This approach computes the covariance matrices with the empirical two-point correlation function, accurate survey geometry and selection effects, but without the contributions of three-point and connected four-point functions.⁸ These higher-point contributions to the large-scale covariance matrix appear similar to the effects of shot noise. As a result, rescaling the shot noise allows to mimic the omitted terms. The amount of rescaling is calibrated on the jackknife covariance matrix estimate from data. The up-to-date methodology was comprehensively described in [97] along with its validation on DESI DR1 mocks. The only new challenge for DR2 covariance estimation we identified is the denser BGS sample (which is also harder to model in mocks), which took more time and displayed slightly worse intrinsic precision. In many other aspects (e.g., footprint boundaries, holes and completeness uniformity) the data has become simpler and more regular than DR1. The code to generate the final DESI DR2 covariances for the correlation functions is accessible on GitHub.⁹

⁷<https://github.com/oliverphilcox/RascalC>.

⁸Only the two-point function and the disconnected four-point function are included.

⁹<https://github.com/cosmodesi/RascalC-scripts/tree/DESI-DR2-BAO/DESI/Y3>, post and pre directories for after and before reconstruction respectively.

TABLE IV. The free parameters and their priors for FS and CS analyses. $\mathcal{N}(\mu, \sigma)$ refers to a normal distribution of mean μ and standard deviation σ , and $[x_1, x_2]$ to a flat distribution between x_1 and x_2 . Parameters marked with * are fixed to the following values when only a 1D fit is performed: $\alpha_{\text{AP}} = 1$, $d\beta = 1$, $a_{2,n} = 0$, $b_{2,n} = 0$.

Parameter	$P(k)$ prior	$\xi(r)$ prior	Description
α_{iso}	[0.8, 1.2]	[0.8, 1.2]	Isotropic BAO dilation
α_{AP}^*	[0.8, 1.2]	[0.8, 1.2]	Anisotropic (AP) BAO dilation
$\Sigma_{\perp}$	$\mathcal{N}(\Sigma_{\perp}^{\text{fid}}, 1.0)$	$\mathcal{N}(\Sigma_{\perp}^{\text{fid}}, 1.0)$	Transverse BAO damping [h^{-1} Mpc]
$\Sigma_{\parallel}$	$\mathcal{N}(\Sigma_{\parallel}^{\text{fid}}, 2.0)$	$\mathcal{N}(\Sigma_{\parallel}^{\text{fid}}, 2.0)$	Line-of-sight BAO damping [h^{-1} Mpc]
Σ_s	$\mathcal{N}(2.0, 2.0)$	$\mathcal{N}(2.0, 2.0)$	FoG damping [h^{-1} Mpc]
b_1	[0.2, 4]	[0.2, 4]	Linear galaxy bias
$d\beta^*$	[0.7, 1.3]	[0.7, 1.3]	Linear RSD parameter
$a_{0,n}$	$\mathcal{N}(0, 10^4)$	Not applicable	Spline parameters for the monopole
$a_{2,n}^*$	$\mathcal{N}(0, 10^4)$	$\mathcal{N}(0, 10^4)$	Spline parameters for the quadrupole
$b_{0,n}$	Not applicable	$[-\infty, \infty]$	Unknown large-scale systematics
$b_{2,n}^*$	Not applicable	$[-\infty, \infty]$	Unknown large-scale systematics
Fitting range	$[0.02, 0.3] h \text{ Mpc}^{-1}$	$[58, 152] h^{-1} \text{ Mpc}$	Measurement bin edges
Data binning	$0.005 h \text{ Mpc}^{-1}$	$4 h^{-1} \text{ Mpc}$	Measurement bin width

For the analyses in FS, we use analytical covariance matrices computed with TheCov¹⁰ [98,99]. The description of the methodology and a comparison with DESI DR1 mock-based sample covariances were presented in [99]. For the purposes of this work, we only compute the disconnected part of the four-point function (sometimes referred to as the Gaussian covariance), including survey geometry effects, and we use the power spectrum measured directly from the corresponding data set as input. Although the connected term for BAO-reconstructed power spectrum multipoles has been recently studied using perturbation theory [100], it has little impact on BAO fits [99] and is therefore neglected in the context of the validation tests presented here.

E. Parameter inference

By combining all components from the previous subsections, the BAO model systematically accounts for the choice of fiducial cosmology, broadband contamination, anisotropies, and nonlinear effects, ensuring accurate recovery of the BAO scale. This approach allows the oscillatory features to serve as a robust standard ruler for cosmological distance measurements.

Unless otherwise noted, our baseline data vector consists of the post-reconstruction monopole and quadrupole moments of the galaxy correlation function (LRGs, ELGs and QSOs), or simply the post-reconstruction monopole in the case of the BGS isotropic BAO fits. A summary of our baseline scale cuts and free parameters of the BAO model is shown in Table IV, along with the prior distributions used during parameter inference.

We sample the posterior distribution with the desilike¹¹ framework, using a wrapper around the Markov chain

Monte Carlo code EMCEE [101]. We assume a Gaussian likelihood and adopt the Gelman-Rubin statistic [102] as the convergence criteria for our chains, demanding that $R - 1 < 0.01$. We also perform maximization using the Minuit profiler [103].

V. KEY CHANGES IN THE DESI DR2 ANALYSIS

This section summarizes the key updates made in the DESI DR2 analysis compared with DESI DR1, highlighting improvements in data handling, methodology, and systematic error treatment. While the systematic error budget is largely inherited from the DESI DR1 BAO analysis, the DESI DR2 analysis introduces tracer-dependent refinements to improve accuracy and robustness. These updates leverage the larger data set and refined analysis techniques to achieve more precise BAO measurements.

A. Data and methodology updates

- (1) *Magnitude cut for BGS*: In the DESI DR1 analysis, the baseline BGS was defined using a magnitude cut of $M_r < -21.5$ (denoted as BGS-BRIGHT-21.5). For DESI DR2, a slightly fainter magnitude cut of $M_r < -21.35$ (BGS-BRIGHT-21.35) was adopted as the baseline. This lower threshold maintains an approximately constant number density to $z < 0.4$, increasing it from $\sim 0.0005 h^3 \text{ Mpc}^{-3}$ to $\sim 0.001 h^3 \text{ Mpc}^{-3}$ after applying completeness corrections (see Fig. 3 in [37]). The effective volume calculation increases by 17% compared with the fiducial -21.5 cut (holding P_0 fixed), without imparting any additional complexity in modeling the sample's clustering or the covariance of its clustering measurements.

¹⁰<https://github.com/cosmodesi/thecov>.

¹¹<https://github.com/cosmodesi/desilike/>.

- (2) *Combining LRG and ELG data*: Unblinding tests now include the combined sample in addition to individual tracers, ensuring robustness across all redshift bins [54,104]. The combined sample is used as the default for cosmological inference.
- (3) *Minimum scale cut ($s_{\min}$)*: For the DESI DR1 analysis, the minimum scale used in the BAO fits was $s_{\min} = 50 h^{-1}$ Mpc. The effect of changing $s_{\min}$ on the quality of fits to mocks was studied extensively by [83], who showed that the recovered α values and their errors are very stable against changes to the minimum scale in the range $50 \lesssim s_{\min} \lesssim 80 h^{-1}$ Mpc. Fits to the blinded DESI DR2 data revealed somewhat large χ^2 values—exceeding the upper boundary set by the mock distributions—when using $s_{\min} = 50 h^{-1}$ Mpc for some redshift bins, which improved significantly when changing the scale cuts to $s_{\min} = 60 h^{-1}$ Mpc. In both cases of $s_{\min}$, the impact on the BAO measurements was negligible. This may suggest that the flexibility of our broadband nuisance parameters is no longer sufficient for the range $50 < s < 60 h^{-1}$ Mpc for some tracers, given the increased S/N of DESI DR2. While this warrants further investigation, the stability of the BAO measurements justifies adopting a more conservative choice of $s_{\min} = 60 h^{-1}$ Mpc for DR2, ensuring that we robustly isolate the s range of the BAO feature.
- (4) *2D BAO fits*: For DESI DR2, 2D BAO fits are employed for ELG1 and QSOs, where the increased S/N ratio in the clustering allows for stable anisotropic BAO measurements, providing additional cosmological information. In these cases, both the monopole and quadrupole moments of the correlation function are fitted simultaneously to constrain α_{iso} and α_{AP} . However, for tracers where the quadrupole S/N is lower, robust determination of α_{AP} becomes more challenging. In these cases, we perform a 1D fit using only the monopole, which primarily constrains α_{iso} , to avoid introducing a weak, non-Gaussian constraint on α_{AP} . This is the case for BGS, where weaker AP distortions at low redshift and higher correlations between α_{iso} and α_{AP} lead us to conservatively adopt a 1D fit as the default. A detailed discussion on the criteria used to assess 1D vs 2D fits, along with supporting tests on mocks and data, is provided in Appendix B.
- (5) *Split tests*: Additional data-splitting tests have been introduced to further validate the robustness of BAO measurements. These include tests based on imaging survey regions to assess the impact of residual imaging systematics, ensuring that variations in survey depth and observational conditions do not bias the results. Additionally, mass- and magnitude-split tests are performed to verify the stability of

BAO measurements across different galaxy populations, probing potential dependencies on intrinsic tracer properties.

B. Systematic error treatment

While the core methodology of the analysis remains consistent with DESI DR1, the updates implemented for DESI DR2 reflect targeted improvements to better account for the increased data volume and enhanced precision. By incorporating tracer-specific and redshift-dependent systematic error treatments, adjusting scale cuts, and adopting conservative fiducial cosmology assumptions, the DESI DR2 analysis ensures robust and unbiased BAO measurements. This section provides a detailed summary of the systematic error contributions, highlighting the careful refinements applied to each tracer in this updated analysis.

- (1) *Fiducial cosmology systematics*: This source of uncertainty quantifies the impact of changing the fiducial cosmology assumed in the BAO fitting pipeline—specifically, the cosmology used to convert redshifts into comoving distances, to evaluate the model templates, and to perform density field reconstruction. To assess this, we rerun the full analysis using an alternative fiducial model and interpret any significant shift in the recovered α_{iso} and α_{AP} values (exceeding a 3σ threshold) as a systematic error [89].

For DESI DR2, this test was updated to use a DESI-motivated evolving dark energy model ($w_0 w_a$ CDM), based on the best-fit results from DR1 BAO + CMB + SN data [29]. This led to an increase in the estimated systematic error on α_{AP} for the LRG sample, from 0.1% in DESI DR1 to 0.18% in DESI DR2. The combined LRG3 + ELG1 sample inherits this increase, as its systematic budget reflects the largest tracer-level contributions. Overall, the total systematic error on α_{AP} increased from 0.3% to 0.335%.

We emphasize that this category of systematics is distinct from what we refer to as “*theory and modeling systematic error*,” which encompasses uncertainties arising from assumptions internal to the BAO modeling framework. These include, for instance, the impact of the dewiggle scheme on BAO extraction, nonlinear mode coupling, relative velocity effects, and choices in broadband modeling. A more complete breakdown of these theoretical contributions can be found in Table 1 of [83].

- (2) *HOD-related systematics*: These systematics quantify the potential bias in BAO parameter recovery due to uncertainty in the galaxy-halo connection. We model this using HOD prescriptions, and define the systematic shift as the change in α_{iso} or α_{AP} when varying HOD parameters across a range of viable models that are consistent with the observed

TABLE V. Systematic error contributions to α_{iso} and α_{AP} for the different DESI tracers. The sources of systematic errors include theoretical systematics, HOD-related systematics, and uncertainties from the fiducial cosmology. The total systematic error for each parameter is obtained by taking the quadratic sum of the individual contributions.

Tracer	Parameter	Theory (%)	HOD (%)	Fiducial (%)	Total (%)
BGS	α_{iso}	0.1	No detection	0.1	0.141
LRG1	α_{iso}	0.1	No detection	0.1	0.141
	α_{AP}	0.2	0.19	0.18	0.329
LRG2	α_{iso}	0.1	No detection	0.1	0.141
	α_{AP}	0.2	0.19	0.18	0.329
LRG2h	α_{iso}	0.1	0.17	0.1	0.221
	α_{AP}	0.2	0.19	0.18	0.329
LRG3 + ELG1	α_{iso}	0.1	0.17	0.1	0.221
	α_{AP}	0.2	0.19	0.18	0.329
ELG1	α_{iso}	0.1	0.17	0.1	0.221
	α_{AP}	0.2	No detection	0.1	0.224
ELG2	α_{iso}	0.1	0.17	0.1	0.221
	α_{AP}	0.2	No detection	0.1	0.224
QSOs	α_{iso}	0.1	0.17	0.1	0.221
	α_{AP}	0.2	0.19	0.18	0.329

data. Any significant difference (exceeding a 3σ threshold) between the recovered values is interpreted as a systematic error from HOD modeling assumptions [88].

For DESI DR1, although tracer-dependent HOD systematics were determined [29,87,88], a conservative approach was taken by adopting the largest detected systematic shift across all tracers as a universal HOD systematic error. For DESI DR2, we refine this treatment by incorporating tracer dependency, applying different HOD systematic contributions for each tracer type. We refer to Table V for a detailed quantification of these values. The updated treatment is as follows:

- (a) *BGS*: No HOD-related systematic error is added, as none were detected within the statistical precision of our test in DESI DR1.
- (b) *LRG1*; *LRG2*: The DESI DR1 values are used, with no systematic error in α_{iso} and a 0.19% error in α_{AP} .
- (c) *ELG1*; *ELG2*: ELG-specific HOD errors are applied, with a 0.17% error in α_{iso} and none detected within the statistical precision of our test in α_{AP} .
- (d) *Combined LRG3 + ELG1*: The largest of the LRG and ELG errors is adopted (0.17% for α_{iso} and 0.19% for α_{AP}).
- (e) *LRG3*: *LRG2h* is treated similarly to the *LRG3 + ELG1* sample, as the redshift distribution $n(z)$ varies significantly across this range, likely accompanied by an evolving bias. Consequently, applying the same HOD systematics as the other LRG bins would likely underestimate the systematic uncertainties for this redshift bin.

- (f) *QSO*: In DESI DR1, no HOD systematics were detected on α_{iso} in the 1D BAO analysis for QSOs. However, for DESI DR2, QSOs were upgraded to a 2D BAO analysis. Since we do not have a pre-determined 2D BAO HOD systematic error for this tracer, we adopt a conservative approach by assigning the same systematic error as the *LRG3 + ELG1* combined sample—not just for HOD, but for all systematic contributions. Given the large statistical uncertainty in QSO measurements, the impact of these systematics is expected to be negligible.

These refinements ensure a tracer-dependent treatment of HOD-related systematics, moving beyond the conservative universal approach of DESI DR1 while maintaining robustness in DESI DR2.

Several systematic effects identified and tested in DESI DR1 were not reassessed in DESI DR2 because they were previously found to be negligible or robust at a precision sufficient for DR2, with no new evidence warranting reevaluation. For example, the theoretical systematics included in Table V remain unchanged from DESI DR1, as their impact was thoroughly evaluated in [83] with a precision of 0.01–0.1%, and there is no reason to expect any change in theoretical systematics since we use the same BAO fitting model and method described in [83]. Similarly, systematic uncertainties related to reconstruction algorithms were studied in [84] using the Y5 footprint, which found no significant contribution to BAO measurements. While the Y3 footprint has a more fragmented and irregular boundary than Y5, we have explicitly verified that this does not introduce bias by analyzing 25 Y3 post-reconstruction AbacusSummit mocks. As shown in Table VII, the recovered BAO parameters remain unbiased across all tracers, including BGS, which is most sensitive to boundary

effects. These findings are consistent with Table 12 of [27], which summarizes the systematic effects considered in DESI DR1.

Although there is no evidence requiring reevaluation, as a sanity check, we allow for theoretical systematics to be correlated across different tracers and redshift bins (see Sec. VII for details).

Additional systematic effects, such as fiber assignment and spectroscopic efficiency corrections, were also not explicitly revisited in detail, as they are indirectly tested through our standard validation procedures. The most stringent test of fiber assignment systematics comes from the requirement that BAO fits yield unbiased results when applied to realistic mock catalogs, as presented in Table VI. These tests confirmed that fiber assignment does not introduce significant bias in BAO measurements, validating the robustness of our correction methods. Furthermore, spectroscopic completeness effects were analyzed in DESI DR1 [105,106] and found to have no measurable impact on BAO constraints. Specifically, [105] examined potential systematics, while [106] studied the impact of observational variations on clustering measurements. These DR1 findings provided sufficient evidence that spectroscopic completeness does not bias BAO measurements, so no additional reanalysis was conducted for DESI DR2.

The calibration accuracy of DESI redshifts has been evaluated to be better than 1 km/s in [107]. We note, however, that the DESI redshift pipeline [45] obtains results in the frame of the solar barycenter. From the CMB dipole, it is known that there is an extra relative velocity of up to 369.82 km/s (see [108] Table 3) that should systematically vary the redshift determination for any objects not bound to the Milky Way. Previous BAO analyses have not accounted for this, and the DESI LSS catalogs used in this study did not include variations in the redshift determination.¹² To check for any potential impact on DESI DR2 BAO measurements, we used the measured direction and amplitude of the CMB dipole to determine the mean redshift offset in each of our DESI DR2 samples. The relative offset is greatest at low redshift, and we find a mean difference of 0.0002 in redshift for the BGS sample. This would result in a negligible correction to any distances.

To summarize, the DESI DR2 analysis builds upon the solid foundation established in DESI DR1. This ensures that our BAO constraints remain robust while leveraging the expanded data set and increased statistical precision of DR2.

VI. RESULTS

As part of the validation process, we conducted our analysis exclusively on blinded catalogs to prevent

¹²The DESI team is currently implementing and testing the inclusion of a correction to redshifts for use in future LSS catalogs.

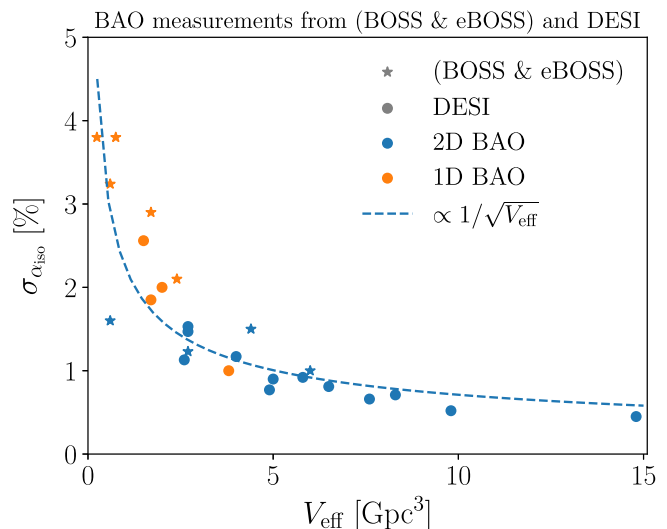


FIG. 3. Comparison of BAO measurement precision as a function of effective survey volume, V_{eff} , for BOSS & eBOSS, including points from [14,109] and DESI [27,37]. The plot illustrates the fractional uncertainty on the BAO scale, σ_{iso} , highlighting the expected scaling relation $\sigma_{\text{iso}} \propto 1/\sqrt{V_{\text{eff}}}$ (dashed line). The DESI data points demonstrate a significant reduction in uncertainty compared with BOSS & eBOSS, benefiting from increased survey volume and improved analysis techniques. The distinction between 1D and 2D BAO fits is also shown, providing historical context for the evolution of BAO precision.

potential confirmation biases. *All statistical checks in this section were initially produced using the blinded data to verify that they met the blinding criteria outlined in Sec. VIB.* Once the blinding tests were passed, we regenerated the plots using the final unblinded catalogs. Therefore, the figures presented in this section now reflect the fully validated, unblinded measurements.¹³

A. Evolution of BAO precision: BOSS & eBOSS to DESI

The precision of BAO measurements has steadily improved with increasing survey volume, advances in data analysis techniques, and improvements in observational hardware. Figure 3 illustrates the relationship between the effective survey volume, V_{eff} , and the fractional uncertainty on the BAO scale, σ_{iso} , for both BOSS & eBOSS and DESI.

The trend shown in the figure follows the expected scaling of $\sigma_{\text{iso}} \propto 1/\sqrt{V_{\text{eff}}}$, highlighting how larger data sets enable more precise BAO constraints. While early BAO measurements, such as those from BOSS & eBOSS, were limited by statistical uncertainties due to smaller galaxy

¹³In the following Sec. VII, we present what we refer to as post-unblinding tests. These tests were planned before unblinding but were only carried out afterward. While they do not influence our primary analysis choices, they provide additional insights into the stability of our results under different assumptions.

TABLE VI. Unblinding checklist. Each test evaluates whether differences observed in the blinded data are consistent with the range spanned by 25 Abacus-2 DR2 mocks. A test is considered successfully passed if the observed differences fall within this range.

No.	Test	Result
1	Are χ^2 reasonable and consistent with mocks?	<i>Yes.</i> Reasonable χ^2 for all tracers, and consistency between data and mocks in all cases.
2	Are the reconstruction settings appropriate for DESI DR2?	<i>Yes.</i> Good performance of reconstruction for all tracers, and data errors consistent with the mocks.
3	Are results robust to imaging systematics?	<i>Yes.</i> BAO constraints remain stable even in the extreme case removing imaging systematics weights from clustering measurements.
4	Are results robust to data splits?	<i>Yes.</i> BAO constraints show robustness when testing different sky regions and magnitude/mass splits.
5	Are uncertainties consistent with mocks?	<i>Yes.</i> Errors on the α parameters measured from data fall within the distribution spanned by the mocks for all tracers.
6	Are the α parameters consistent between pre- and post-reconstruction?	<i>Yes.</i> Pre-reconstruction α_{iso} values show slight deviations relative to mocks but remain within expected variations. Reconstruction improves precision, and the shifts in α values are consistent with the mock distribution.
7	Are mock fits unbiased?	<i>Yes.</i> Most tracers show biases below 0.5σ . The largest offset is 2.46σ for post-reconstruction QSOs, which remains below our 3σ threshold for systematic detection.
8	Are α values consistent between CS and FS?	<i>Yes.</i> Results from the correlation function are consistent to within 0.5σ with those from the power spectrum.
9	Are we reassessing the systematic error budget?	The systematic error budget is primarily based on DESI DR1, with updates as described in Sec. VB.
10	Are LRGs and ELGs consistent in the overlapping bin $0.8 < z < 1.1$?	<i>Yes.</i> The α_{iso} offset between LRG2h and ELG1 is 0.41σ , consistent with mock expectations. The agreement on α_{AP} is within -1.66σ .

samples, DESI has significantly expanded the survey volume, leading to a substantial reduction in σ_{iso} . While DESI remains statistically limited, the significant increase in data set size has driven a major improvement in measurement precision, complemented by refined analysis techniques.

Additionally, the figure differentiates between 1D and 2D BAO analyses. While 2D BAO fits extract additional information from anisotropic clustering, 1D fits provide robust constraints in lower S/N regimes, such as for BGS in DESI (see Appendix B for details). This historical comparison contextualizes the improvements made in BAO analyses and demonstrates DESI’s capability to push the limits of precision cosmology through larger survey volumes and improved methodologies.

B. The unblinding tests

To ensure the robustness of our DESI DR2 BAO measurements, we conducted a series of pre-unblinding validation checks, which are summarized in Table VI. These tests verify the consistency of the clustering measurements across different tracers, data splits, and methodologies while assessing systematic uncertainties. Each test is evaluated by comparing the observed differences in the blinded data to the full range spanned by the 25 Abacus-2 DR2 mocks. A test is considered successfully passed if the

observed differences fall within this range. We determined the systematic error budget detailed in Table V prior to unblinding.

Following DR1, we define a systematic effect as one that exceeds a significance of 3σ , where σ refers to the statistical precision associated with the given test.

When introducing our various results in the following subsections, we will discuss how each of them helps to assess one or more items from the unblinding checklist (Table VI).

C. Consistency and precision in DESI DR1 and DESI DR2 two-point clustering measurements

We compare the two-point clustering measurements from DESI DR1 and DESI DR2, as shown in Fig. 4.¹⁴ This comparison serves two purposes: assessing the consistency between data releases with different footprints and completeness and highlighting improvements in measurement precision. An analogous plot (DESI DR1 vs DESI DR2) but for the mean of the Abacus-2 mocks is shown in Appendix C.

¹⁴We emphasize that the DR1–DR2 comparison shown in Fig. 4 was created only after the DR2 results were validated and unblinded. No such comparison was examined during the blinded phase, and it did not inform any analysis decisions or interpretations prior to unblinding.

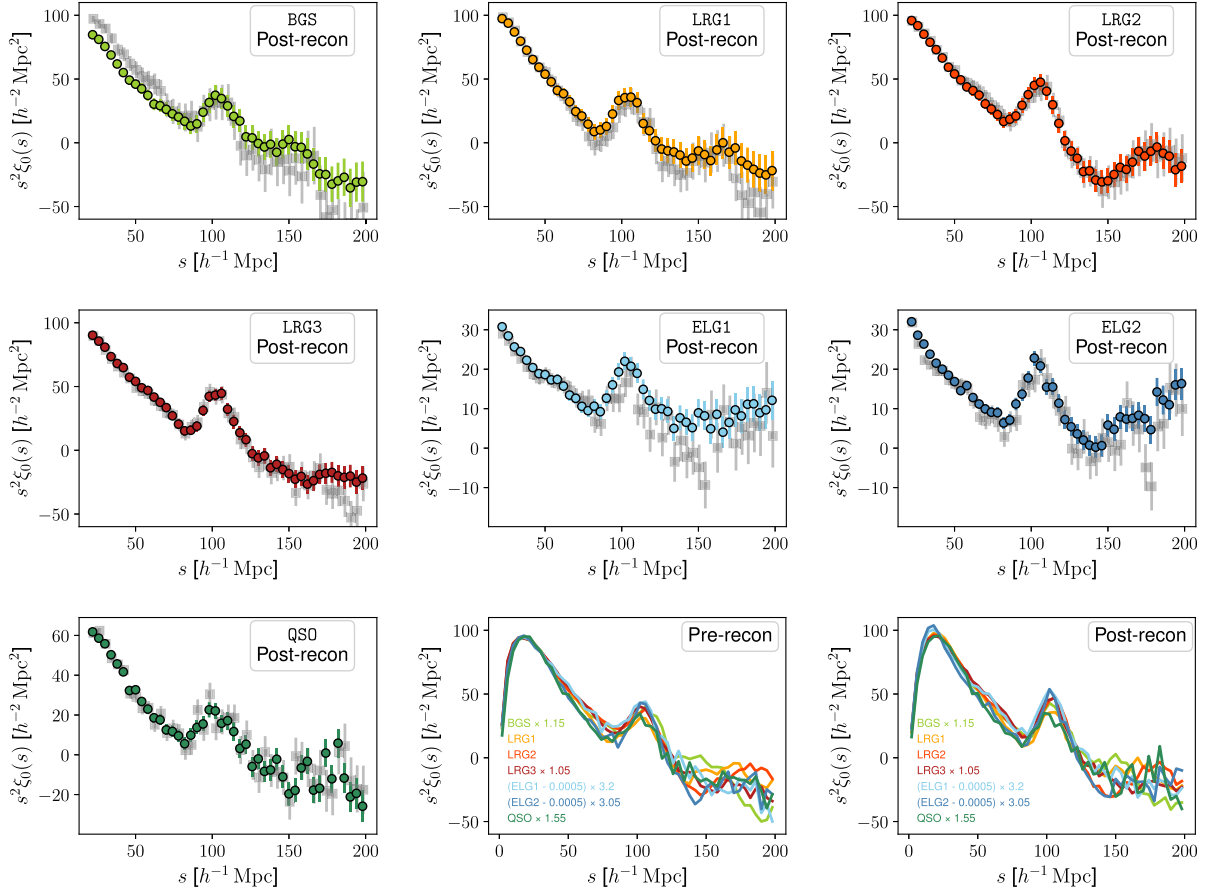


FIG. 4. Comparison of the monopole of the unblinded two-point correlation function, $\xi_0(s)$, between DESI DR1 (gray) and DESI DR2 (colored) for various tracer samples and redshift ranges. The panels show results for BGS ($0.1 < z < 0.4$), LRGs ($0.4 < z < 0.6$, $0.6 < z < 0.8$, $0.8 < z < 1.1$), ELGs ($0.8 < z < 1.1$, $1.1 < z < 1.6$), and QSOs ($0.8 < z < 2.1$). The bottom-middle and bottom-right panels show combined views of all tracers before (Pre-recon) and after (Post-recon) BAO reconstruction, respectively, highlighting the sharpening of the BAO feature after reconstruction. The DR2 measurements exhibit reduced scatter, particularly for the BGS and ELG samples, reflecting improved precision. Notable differences in amplitude, especially for BGS at small scales and ELGs at large scales, are discussed in the text and are attributed to changes in sample selection and potential residual imaging systematics, respectively.

Across all samples, the DR2 data are noticeably smoother, indicating enhanced precision in the clustering measurements. This is particularly evident in the ELG1 and BGS samples, where the BAO feature appears sharper in DR2. The clustering amplitudes between DR1 and DR2 are generally consistent, suggesting that the sample properties have remained stable across the survey footprint. However, the BGS sample presents a notable exception, with a shift in the small-scale amplitude. This difference arises from the application of a fainter absolute magnitude cut in the DR2 sample (see Sec. V), which affects the clustering signal at smaller scales. For the ELG samples—especially ELG1—we observe higher large-scale amplitudes in DR2, which could indicate residual imaging systematics. These two effects are mitigated as follows.

The bottom-middle and bottom-right panels of Fig. 4 illustrate the impact of potential residual systematics in DR2, showing the clustering measurements across all tracers before (pre-reconstruction) and after (post-reconstruction)

BAO reconstruction, respectively. The observed differences between DR1 and DR2 clustering amplitudes can largely be accounted for by an approximate multiplicative factor of 1.15 for BGS and a constant offset of 0.0005 for ELGs. Importantly, our BAO fitting model includes both a bias multiplicative term and a free constant offset, ensuring that these residual effects do not significantly impact the BAO results. Furthermore, [55] demonstrated that DESI BAO measurements from the ELG sample in DR1 remained robust even in the absence of explicit imaging systematics corrections. Since the level of residual contamination in DR2 remains within a comparable range to that in DR1, it does not pose a concern for the BAO measurements presented here.¹⁵

¹⁵Future work will investigate these residuals further, particularly for models that rely on the broadband clustering amplitude, where such contamination may have a larger effect.

D. BAO measurements from DESI DR2

We now focus on the BAO constraints derived from the CS measurements shown in the previous section and compare them with the expectations from the mocks.¹⁶ For the resulting BAO constraints on the dilatation parameters $\{\alpha_{\text{iso}}, \alpha_{\text{AP}}\}$ and a visualization of the BAO features on the unblinded data, see Table III and Fig. 5 of the key paper [37]. To assess the robustness of our BAO measurements, we compare our results against both pre-reconstruction and post-reconstruction fits, validate the uncertainty estimates using mocks, and evaluate statistical goodness-of-fit metrics. This section presents key tests addressing the consistency of our results.

1. Pre- vs post-reconstruction BAO fits

Figure 5 compares the pre- and post-reconstruction BAO fits for the DESI DR2 data and the 25 Abacus-2 DR2 mocks. With the exception of the first panel, which shows the BGS tracer and employs 1D (α_{iso}) fits, all other tracers use 2D($\alpha_{\text{iso}}, \alpha_{\text{AP}}$) fits. For the 1D BAO fits, we use only the monopole of the correlation function. The data measurements (indicated by stars) fall comfortably within the range spanned by the 25 Abacus-2 DR2 mocks. This range serves as our benchmark for defining consistency, given the limited number of mock realizations.

Reconstruction improves the precision on α_{iso} for all tracers. For DR2, the percentage of improvement compared with the pre-reconstruction fits ranges from 3.7% for QSOs, to 42% for LRG3 + ELG1, varying depending on the signal to noise of the density field and also the severity of degradation to mitigate. α_{AP} also benefits greatly from reconstruction, with a gain that ranges from 8.5% for QSOs, to 47% for LRG3 + ELG1. This level of improvement from the data finds good agreement with the distribution of mocks and is also consistent with the DR1 findings, emphasizing the robustness of the reconstruction pipeline. A point to be noted is that in DR1, the QSOs saw a small degradation of the constraining power on α_{iso} after reconstruction, which we attributed to a noise fluctuation in the data since this was still consistent with the range of possibilities spanned by the mock catalogs. This is confirmed by the new DR2 results, where the QSOs sample shows a small improvement from reconstruction, while now also allowing for a separate constraint on α_{AP} due to the promotion of this tracer to a 2D fit. Overall, these results address the unblinding test no. 2, showing that the adopted reconstruction settings lead to the expected improvements with respect to the pre-reconstruction constraints. Figure 5 also addresses the unblinding test no. 6, showing consistency between the best-fit values of the BAO scaling

parameters before and after reconstruction. Finally, it also informs unblinding test no. 5, demonstrating that the uncertainties on the BAO scaling parameters, both before and after reconstruction, are consistent with the distribution of mocks. This confirms that the relative improvement in individual errors aligns with the average expected trend.

2. Robustness of the DESI DR2 uncertainties

Figure 6 provides a visual check on our uncertainty estimates by comparing the normalized deviations $(\alpha_{\text{iso}} - \langle \alpha_{\text{iso}} \rangle) / \sigma_{\alpha_{\text{iso}}}$ and $(\alpha_{\text{AP}} - \langle \alpha_{\text{AP}} \rangle) / \sigma_{\alpha_{\text{AP}}}$ from individual realizations (scatter points) to the expected statistical scatter from a multivariate Gaussian. The dashed circles indicate the 1σ (68%) and 2σ (95%) confidence regions for a multivariate Gaussian distribution centered at (0, 0). If the uncertainties are well characterized, we expect the individual realizations to follow the statistical properties of such multivariate Gaussian contours.

Across all tracers and redshift bins, the scatter of values generally aligns with the expected contours. We expect one mock realization to lie outside the 2σ contour and eight outside the 1σ contour. The obtained results vary from 1 to 3 and 5 to 10, respectively. The error estimates are thus broadly consistent with statistical expectations. This serves as a complementary consistency check for unblinding test no. 5, suggesting that our covariance matrix provides a reasonable description of the statistical scatter within the available set of mock realizations. The limited number of mock realizations (25) and the replication involved in their creation (see Sec. III) prevents this test from being more precise. The most thorough evaluation of the accuracy of uncertainty on BAO measurements recovered when applying *RascalC* covariance matrices remains the DR1 studies [97], which imply an accuracy of better than 5%.

3. Consistency of LRG and ELG samples

All DESI tracers are split by type and redshift when analyzing subsamples in [37], with the exception that we use a combination of LRG3 + ELG1 in the redshift range $0.8 < z < 1.1$ [54,104]. In Fig. 7, we have calculated the offset in α constraints between LRG2h and ELG1 from DR2 data. We find these to be $\Delta\alpha_{\text{iso}}(\text{LRG2h} - \text{ELG1}) = 0.41\sigma$ and $\Delta\alpha_{\text{AP}}(\text{LRG2h} - \text{ELG1}) = -1.66\sigma$. We have explicitly verified that these offsets are compatible with noise fluctuations by comparing them to the same output generated from the 25 mock realizations. This serves as a validation of the consistency of the LRG2h and ELG1 samples (unblinding test no. 10), enabling their combination into the LRG3 + ELG1 tracer that is used for the cosmological interpretation. The covariance for the combined tracer was computed separately with the *RascalC* code similarly to other (single) tracers using the combined catalogs and the corresponding correlation function

¹⁶Additionally, FS results, detailed in Appendix A, further support our finding in CS, given the level of agreement between the two conjugate spaces.

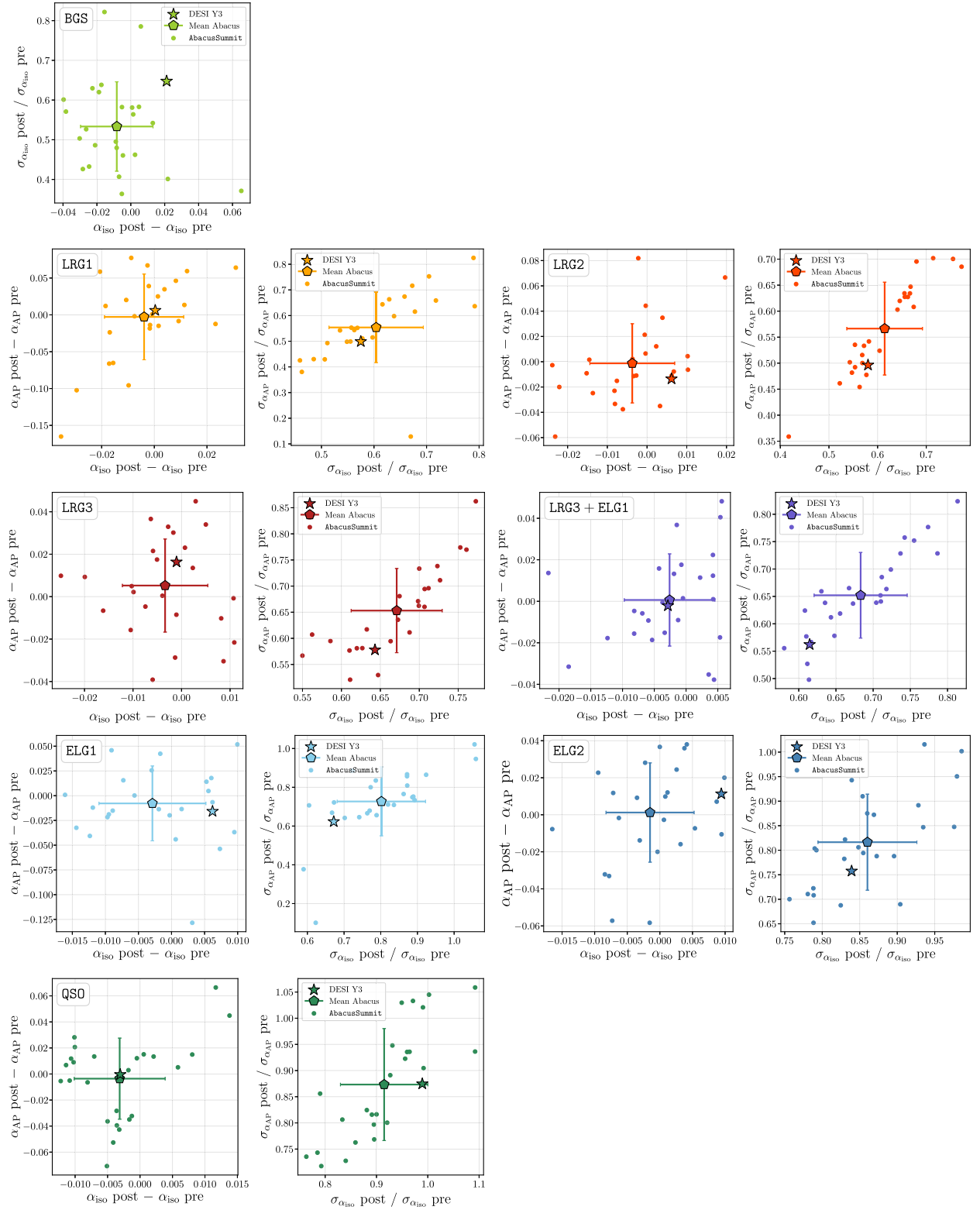


FIG. 5. Pre- and post-reconstruction BAO measurements, comparing DESI DR2 measurements (stars) with the 25 Abacus-2 DR2 (the open points with the error bars are the means and the standard deviations around the means of the 25 mocks). All DESI DR2 fits are consistent with what is expected from the mocks.

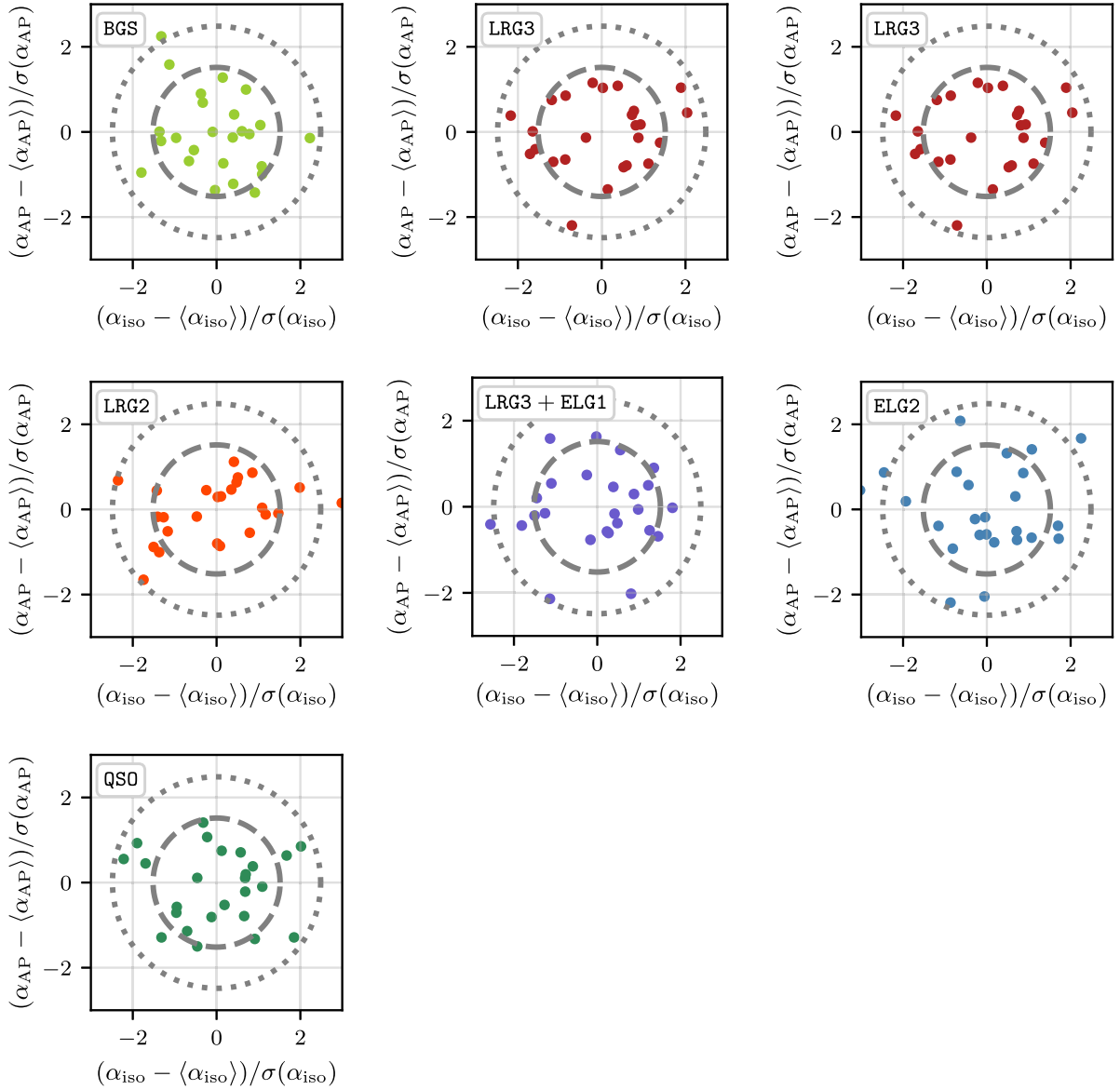


FIG. 6. Comparison of the estimated uncertainties $\sigma_{\alpha_{\text{iso}}}$ and $\sigma_{\alpha_{\text{AP}}}$ across the 25 Abacus-2 DR2 mocks. Scatter points represent individual best-fit realizations, while the dashed circles indicate the expected 1σ (68%) and 2σ (95%) confidence regions of a multivariate Gaussian distribution centered at zero. Each panel corresponds to a different tracer and redshift bin, illustrating the relationship between individual mock uncertainties and the expected statistical spread. The general agreement between the scatter points and the contours indicates that our uncertainty estimates are robust and do not show significant biases.

measurements (including jackknife) [54]. Thus, it takes the LRG2h-ELG1 cross-correlations into account.

4. Goodness-of-fit and χ^2 tests

Validation using DR2 BAO measurements. The χ^2 values for the BAO fits, presented in Table VIII, indicate that the models provide statistically reasonable fits to the data. To quantify this, we report the probability to exceed (PTE), which measures the likelihood of obtaining a χ^2 value as large or larger than the observed one, assuming the model is a good fit. As a general guideline, PTE values close to 0 or 1 suggest a poor fit, as they indicate that the observed χ^2 is

either much lower or much higher than expected. Conversely, values near 0.5 indicate a fit consistent with statistical expectations. In our results, the PTE values range from 0.01 (LRG2h) to 0.92 (QSOs). Only LRG2h falls below the conventional 0.05 threshold. While this could raise concerns in isolation, we emphasize that we are testing nine independent samples, and observing a minimum PTE as low as 0.01 is not unexpected in such a context. In fact, the probability that the lowest PTE among nine independent draws exceeds 0.01 is approximately 91%, indicating that this level of variation is well within statistical expectations.

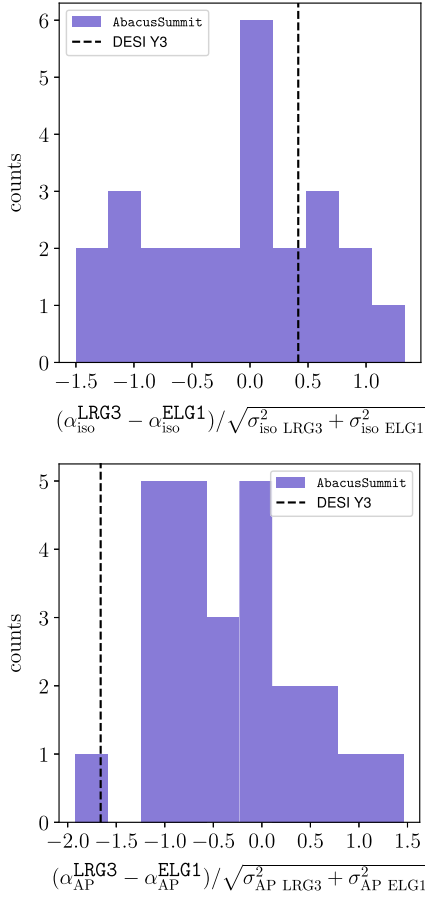


FIG. 7. Offsets in α_{iso} (top) and α_{AP} (bottom) between the LRG2h and ELG1 samples for the DESI DR2 data (black) compared with the distribution from 25 Abacus-2 DR2 mocks (purple). The observed differences, $\Delta\alpha_{\text{iso}} = 0.41\sigma$ and $\Delta\alpha_{\text{AP}} = -1.66\sigma$, are consistent with noise, validating the combination of LRG2h and ELG1 into LRG3 + ELG1 for cosmological analysis.

Additionally, LRG2h is not used in isolation in the cosmological analysis; it is combined with the ELG1 sample to form the LRG3 + ELG1 bin, which has a much more typical PTE of 0.24. The other samples (BGS, LRG1, LRG2, ELG1, ELG2) all yield PTE values between 0.05 and 0.89, consistent with statistical expectations. The QSOs sample shows a high PTE of 0.92, which is consistent with the lower S/N of that measurement and expected statistical fluctuations. Overall, these results confirm that our BAO measurements are statistically robust and that the reconstructed correlation functions provide a reasonable fit to the data. This directly addresses unblinding test no. 1, which is considered satisfactorily passed for all tracers.

Validation using mock catalogs. To further assess the reliability of our BAO error estimates, we compare the measured values to expectations derived from mock catalogs. Table VII presents results from fits to the mean of 25 mock realizations, where the covariance matrix has been rescaled by a factor of 25. This ensures that the errors

reflect the combined volume of all realizations, providing a stringent test of the unbiasedness of the mocks. None of the fits deviate from the expectation of $\alpha_{\text{iso}} = \alpha_{\text{AP}} = 1$ by more than 3σ , reinforcing the reliability of our modeling framework. Pre-reconstruction measurements show a slight tendency for α values larger than one, a well-known effect of nonlinear gravitational evolution, which systematically shifts the BAO position to smaller separations (resulting in larger α values). This bias is significantly reduced after reconstruction, confirming that our pipeline effectively corrects for these nonlinear shifts. These findings address unblinding test no. 7, demonstrating that our mock-based predictions align well with our observed results.

E. Robustness of the DESI DR2 measurements

We have extensively examined the impact of several variations around our baseline settings on the BAO constraints, which are shown in Fig. 8. Results are expressed in terms of the absolute difference in α_{iso} and α_{AP} relative to the baseline, defined as $\Delta\alpha_{\text{iso}} = \alpha_{\text{iso}}^{\text{case}} - \alpha_{\text{iso}}^{\text{baseline}}$ and $\Delta\alpha_{\text{AP}} = \alpha_{\text{AP}}^{\text{case}} - \alpha_{\text{AP}}^{\text{baseline}}$. Each test corresponds to a different robustness check performed to assess the stability of the BAO constraints under various assumptions. To interpret the magnitude of each shift, we also express it in units of the baseline uncertainty—i.e., how many σ the test result deviates from the baseline.

Pre-recon: The first comparison is with the pre-reconstruction correlation function fits. The mean values of the parameters show expected shifts relative to the baseline, with variations that generally do not exceed 1σ , except for the BGS sample, which deviates by 1.18σ . These shifts are consistent with the nonlinear displacement of the BAO feature, which reconstruction is designed to correct. The level of improvement observed post-reconstruction closely matches expectations from DR2 mocks (see Fig. 5). As expected, parameter precision is reduced in the pre-reconstruction fits compared with the baseline, reconstructed measurements. The gain in precision after reconstruction aligns well with the trends observed in the mocks, reinforcing the robustness of our analysis. This comparison further validates unblinding tests no. 2 and no. 6.

1D BAO fit: Next, we examine the constraints obtained using a 1D BAO fit, where the only dilation parameter that is varied is α_{iso} , with $\alpha_{\text{AP}} = 1$ fixed. BGS is the only tracer that uses a 1D fit by default, and therefore, this constraint agrees exactly with the baseline by construction. The α_{iso} constraints from 1D fits show excellent agreement with their 2D counterparts, in terms of both precision and mean values. The deviations from the baseline range from 0.01σ (LRG2h) to 0.62σ (ELG1), fully consistent with statistical fluctuations expected from the mocks.

Power spectrum BAO fit: Constraints from the galaxy power spectrum show good agreement with the CS results when compared with the agreement expected from the

TABLE VII. Constraints on the BAO scaling parameters from fits to the mean of 25 Abacus-2 DR2 mocks, where the average errors have been rescaled by $\sqrt{25}$ so that the errors reflect the combined volume of all mocks. The lower and upper sub-panels show results before and after reconstructing the galaxy catalogs. The fifth and sixth columns show the offset, in terms of the number of standard deviations, from the expectation of $\alpha_{\text{iso}} = \alpha_{\text{AP}} = 1$. r is the cross-correlation coefficient between α_{iso} and α_{AP} . We also show the χ^2 per degree of freedom at the best fit.

Tracer	Reconstruction	α_{iso}	α_{AP}	$\Delta\alpha_{\text{iso}}$	$\Delta\alpha_{\text{AP}}$	r	χ^2/dof
BGS	Post	0.9980 ± 0.0020		-0.97			18.0/16
LRG1	Post	0.9971 ± 0.0017	0.9941 ± 0.0060	-1.70σ	-0.99σ	0.102	30.3/33
LRG2	Post	0.9993 ± 0.0013	1.0012 ± 0.0049	-0.49σ	0.25σ	0.148	24.0/33
LRG3	Post	0.9991 ± 0.0012	0.9956 ± 0.0042	-0.76σ	-1.05σ	0.111	23.2/33
ELG1	Post	1.0009 ± 0.0021	1.0096 ± 0.0070	0.43σ	1.37σ	-0.231	38.1/33
ELG2	Post	1.0004 ± 0.0015	0.9982 ± 0.0054	0.26σ	-0.33σ	-0.135	24.6/33
LRG3 + ELG1	Post	1.0004 ± 0.0010	1.0001 ± 0.0037	0.36σ	0.02σ	0.082	38.8/33
QSOs	Post	0.9943 ± 0.0023	1.0049 ± 0.0091	-2.46σ	0.54σ	-0.022	29.1/33
BGS	Pre	1.0094 ± 0.0041		2.29σ			27.0/16
LRG1	Pre	1.0004 ± 0.0031	0.9898 ± 0.0113	0.14σ	-0.91σ	0.327	38.9/33
LRG2	Pre	1.0041 ± 0.0023	1.0099 ± 0.0094	1.76σ	1.05σ	0.471	18.2/33
LRG3	Pre	1.0030 ± 0.0018	0.9930 ± 0.0067	1.68σ	-1.03σ	0.157	28.3/33
ELG1	Pre	1.0043 ± 0.0026	1.0197 ± 0.0099	1.63σ	1.99σ	-0.107	49.3/33
ELG2	Pre	1.0015 ± 0.0018	0.9951 ± 0.0067	0.81σ	-0.73σ	-0.034	19.9/33
LRG3 + ELG1	Pre	1.0035 ± 0.0015	1.0021 ± 0.0059	2.25σ	0.36σ	0.182	34.9/33
QSOs	Pre	0.9979 ± 0.0025	1.0079 ± 0.0106	-0.83σ	0.74σ	0.097	16.2/33

mocks. This comparison is explored in more detail in Appendix A, where we assess the consistency of power spectrum and correlation function fits under different data vector variations. While these two statistics should, in principle, contain the same amount of information, practical differences arise due to the adopted scale cuts, modeling choices, and particularly the treatment of the covariance matrix. These factors can lead to small variations in the derived constraints. The BGS sample presents the largest deviation, showing a 0.55% offset, which is slightly larger than the most extreme mock realization but remains within an acceptable range. The overall agreement reinforces the robustness of our methodology and ensures that any differences between the two approaches are well

TABLE VIII. Summary of χ^2 goodness-of-fit values for the BAO measurements across different tracers and redshift bins using unblinded DESI DR2 data. The PTE quantifies the likelihood of obtaining a χ^2 value as large or larger than the observed value, assuming the model provides a good fit to the data. This serves as a diagnostic check to assess whether the BAO fits are statistically reasonable.

Tracer	Redshift range	$\chi^2/\text{d.o.f.}$	PTE
BGS	$0.1 < z < 0.4$	26.09/16	0.053
LRG1	$0.4 < z < 0.6$	40.27/33	0.180
LRG2	$0.6 < z < 0.8$	23.44/33	0.891
LRG2h	$0.8 < z < 1.1$	53.77/33	0.013
LRG3 + ELG1	$0.8 < z < 1.1$	38.44/33	0.237
ELG1	$0.8 < z < 1.1$	26.87/33	0.765
ELG2	$1.1 < z < 1.6$	42.39/33	0.127
QSOs	$0.8 < z < 2.1$	22.19/33	0.923

understood. These results directly inform the unblinding test no. 8.

Flat priors on the damping scales: Our baseline configuration assumes Gaussian priors for the BAO damping and FoG parameters [Eq. (15)], informed from fits to mock catalogs [83]. We find that switching these to uniform priors does not significantly affect the constraints on the scaling parameters, with only a very mild degradation in the α_{iso} precision of the QSOs. However, fits on multiple realizations of mock catalogs showed more stability of the BAO fits when adopting an appropriately chosen Gaussian prior, ensuring unbiasedness of the parameter constraints without compromising precision [83].

Polynomial broadband (iso for monopole only): Our parametrization of the broadband shape of the power spectrum involves a new spline basis [Eq. (17)] that was introduced in [83] for the DR1 BAO analysis. The “polynomial broadband” row in Fig. 8 shows constraints obtained using the polynomial-fitting method similar to the approach used in BOSS [110], which are perfectly consistent with the results from the new spline method.

Fixing broadband (iso for monopole only): We test the impact of broadband modeling on our BAO results by comparing cases where the broadband is fixed for both the monopole and quadrupole of the correlation function (“fixed broadband”) versus only for the monopole (“fixed broadband iso”). The results show a very mild impact, with the largest observed shift being 0.55σ in α_{iso} for ELG1.

The effect of the $d\beta$ prior: The $d\beta$ parameter quantifies the ratio between the true value of the growth rate of structures f and that predicted by the fiducial cosmology. The bounds of the uniform prior that is used as our baseline

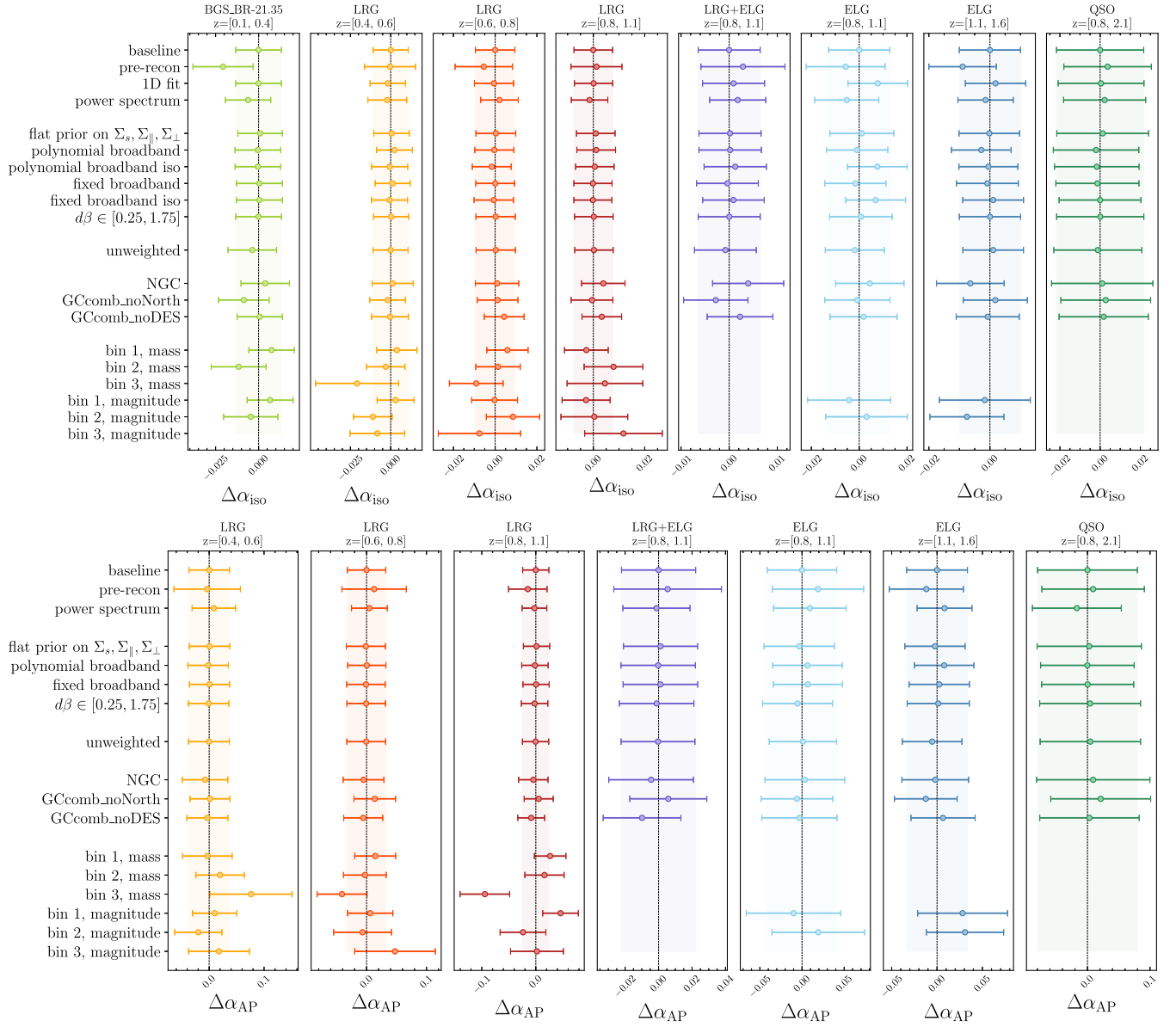


FIG. 8. Response of the constraints on the isotropic (top) and anisotropic (bottom) BAO scaling parameters to changes in the data vector or model assumptions, using unblinded DESI DR2 data. The first row represents the difference between the baseline BAO constraints and itself, resulting in values centered at zero, with the shaded regions indicating the statistical uncertainty of the baseline measurements. The remaining rows show the difference $\Delta\alpha$ between the BAO scaling parameters obtained in each test and the baseline value. Specifically, $\Delta\alpha_{\text{iso}} = \alpha_{\text{iso}}^{\text{case}} - \alpha_{\text{iso}}^{\text{baseline}}$ and $\Delta\alpha_{\text{AP}} = \alpha_{\text{AP}}^{\text{case}} - \alpha_{\text{AP}}^{\text{baseline}}$. Note the distinction in the BAO analysis: for the BGS tracer, 1D fits are used, whereas 2D fits are applied for the other tracers.

are [0.7, 1.3], which we deem a reasonably relaxed prior given the constraining power of our data. We tested relaxing those priors even further to [0.25, 1.75], finding a negligible impact on the BAO constraints.

Unweighted imaging systematics: The correlation between the observed galaxy density and certain imaging properties of the survey in the same sky region produces “imaging systematics” that can bias the clustering measurements. To mitigate them, the LSS catalogs incorporate weights that are added using a regression method that is

informed by the maps of different observational properties. Although the clustering itself can be significantly affected by these weights, the “unweighted” row in Fig. 8 shows that BAO remains robust even when the weights are completely removed from the catalog. For further validation, [55] specifically demonstrated that DESI BAO measurements for ELGs, the sample that is most susceptible to these effects, remain unaffected by imaging systematics. This directly addresses the unblinding test no. 3.

NGC, noNorth, noDES: We perform BAO fits under various data splits, described in Sec. II B. We observe a reduction in precision when restricting the analysis to the NGC region, which is expected given the decreased sky coverage (30% and 44% less area for bright-time and dark-time tracers, respectively). However, parameter shifts remain consistent at the 1σ level. A smaller reduction in precision is observed when excluding the North (31% and 20% of the DR2 sky area for bright and dark time, respectively) or the DES (7% and 9%) imaging regions, with posterior mean shifts remaining within 1σ of the baseline.¹⁷ To further quantify the robustness of these survey splits, we performed a Kolmogorov-Smirnov (KS) test on the significance levels (σ) of the shifts in $(\alpha_{\text{DR2}} - \alpha_{\text{sub}})/\sigma_{\text{diff}}$. The KS test assesses whether a given sample follows an expected probability distribution by comparing its empirical cumulative distribution function (CDF) to a theoretical CDF. In our case, under the null hypothesis, the shifts in $\alpha_{\text{DR2}} - \alpha_{\text{sub}}$ should be Gaussian-distributed noise, implying that the computed σ values should follow a standard normal distribution, $\mathcal{N}(0, 1)$. The KS test quantifies the maximum difference between the empirical CDF of the observed σ values and the CDF of a unit Gaussian, providing a measure of how well the data matches the expected distribution. To construct the test statistic, we computed the observed shifts as $\alpha_{\text{DR2}} - \alpha_{\text{sub}}$, which represent the differences in the BAO scale measurements between the full DR2 sample and each subsample. To evaluate the significance of these shifts, we estimated the corresponding uncertainty, σ_{diff} , directly from the data using

$$\sigma_{\text{diff}}^2 \equiv \sigma_{\text{DR2}}^2 \left(\frac{N_{\text{DR2}}}{N_{\text{sub}}} - 1 \right),$$

where σ_{DR2}^2 is the variance of α_{DR2} , and N_{DR2} and N_{sub} represent the number of galaxies in the full sample and subsample, respectively.¹⁸ This definition ensures that the

¹⁷For further details on the impact of using photometric redshifts from the DESI Legacy Imaging Survey on the BAO analysis, see [111].

¹⁸This formula follows from the standard variance scaling relation for a subsample:

$$\sigma_{\text{sub}}^2 = \sigma_{\text{DR2}}^2 \frac{N_{\text{DR2}}}{N_{\text{sub}}}.$$

Since the subsample is drawn from the full data set, the uncertainty difference is set by the sample size ratio, ensuring that $\sigma_{\text{sub}}^2 - \sigma_{\text{DR2}}^2$ remains positive. This approach properly accounts for the correlation between the full sample and its subsamples, making variance subtraction a valid choice in this specific case. As a result, the uncertainty associated with the shift in α can be expressed as

$$\sigma = \sqrt{\sigma_{\text{sub}}^2 - \sigma_{\text{DR2}}^2},$$

which remains well-defined since σ_{sub}^2 has been properly rescaled. This formulation assumes a *maximum statistical correlation* between the subsample and the full sample, making it a conservative estimate of the uncertainty.

uncertainty accounts for the relative sample sizes and provides a robust estimate of expected fluctuations. By normalizing the observed shifts using σ_{diff} , we obtained a standardized significance measure that allows direct comparison across different subsamples.

In total, we tested 48 cases, incorporating the BGS, LRG1, LRG2, LRG2h, LRG3 + ELG1, ELG1, ELG2, and QSO tracers for both α_{iso} and α_{AP} , comparing results for the full sample against the NGC, GCcomb-noDES, and GCcomb-noNorth regions. We obtained a KS statistic of $D = 0.1489$ with a corresponding p -value of 0.215, indicating no significant deviation from normality. These results suggest that the observed shifts in $\alpha_{\text{DR2}} - \alpha_{\text{sub}}$ are consistent with statistical fluctuations and do not indicate systematic biases across the different survey splits.

Mass and magnitude splits: The final set of tests assesses the robustness of the BAO measurements against variations in tracer properties by performing mass and magnitude splits, as shown in the last six rows of Fig. 8. Since galaxy bias b depends on stellar mass and luminosity, these properties can influence the amplitude and shape of the two-point correlation function; see, for instance, the differences in clustering amplitudes shown in Table IX. To test this, we divide the BGS, LRG, and ELG samples into subsamples based on mass and magnitude, following the methodology outlined in Sec. II. We note that we used the same bias values as in the baseline case, which are depicted in Table II. We merge the mass/magnitude-split

TABLE IX. Summary of clustering amplitudes in mass/magnitude-split subsamples. The clustering amplitudes $\xi_0^{s=20}$ are defined as the two-point correlation function monopole ξ_0 at $s = 20$ Mpc/h. In each block, the amplitude decreases from bin 1 to bin 3. The parameters m_r , m_{W1} , and m_g correspond to the extinction-corrected apparent magnitudes in the r -band, $W1$ -band, and g -band, respectively, as measured in the Legacy Survey DR9. These magnitudes were used to define the magnitude splits in the BGS, LRG, and ELG samples.

Tracer	Magnitude bin	$\xi_0^{s=20}$	Mass bin	$\xi_0^{s=20}$
BGS	m_r bin 1	0.250	M_* bin 1	0.276
BGS	m_r bin 2	0.195	M_* bin 2	0.177
LRG1	m_{W1} bin 1	0.313	M_* bin 1	0.339
LRG1	m_{W1} bin 2	0.244	M_* bin 2	0.259
LRG1	m_{W1} bin 3	0.212	M_* bin 3	0.180
LRG2	m_{W1} bin 1	0.294	M_* bin 1	0.333
LRG2	m_{W1} bin 2	0.241	M_* bin 2	0.254
LRG2	m_{W1} bin 3	0.223	M_* bin 3	0.190
LRG2h	m_{W1} bin 1	0.293	M_* bin 1	0.304
LRG2h	m_{W1} bin 2	0.229	M_* bin 2	0.226
LRG2h	m_{W1} bin 3	0.193	M_* bin 3	0.186
ELG1	m_g bin 1	0.084		
ELG1	m_g bin 2	0.079		
ELG2	m_g bin 1	0.091		
ELG2	m_g bin 2	0.082		

subsamples and perform reconstruction, then extract the shifted subsamples by their mass/magnitude bin labeled before the reconstruction.

Although some deviations from the baseline measurements are visually observed in Fig. 8, a quantitative consistency check confirms that the α_{iso} and α_{AP} values remain statistically consistent with our baseline results. Specifically, we performed a KS test, as detailed previously, on all mass and magnitude subsample cases. The results yielded a p -value of 0.73, indicating that the observed variations across mass- and magnitude-split subsamples are consistent with statistical fluctuations. To quantify the uncertainty, we estimate the joint uncertainty in the KS test as

$$\sigma = \sqrt{\sigma_{\text{sub}}^2 - \sigma_{\text{DR2}}^2},$$

assuming a maximum statistical correlation between the subsamples and the baseline sample, ensuring a conservative uncertainty estimate. The maximum difference is in α_{AP} of mass bin 3 of LRG2h, ranging from 2.0σ to 2.4σ , depending on the value of correlation coefficient. These results confirm that BAO measurements are robust to mass and magnitude selection effects. The combination of all data-split tests informs the unblinding test no. 4, which is considered satisfactorily passed.

Overall, our extensive suite of robustness tests demonstrates the stability of the BAO measurements against a range of potential systematics. We performed tests involving variations in modeling assumptions, data-vector choices, survey region splits, mass/magnitude selections, and imaging systematics. The high level of consistency observed across all tests strongly supports the reliability of our measurements. The agreement across these complementary robustness checks suggests that our constraints are not significantly biased by observational systematics, sample selection, or modeling assumptions. This reinforces the robustness of the DESI DR2 BAO analysis and its suitability for cosmological interpretation.

F. Gaussianity of the BAO posterior

Once the posterior distribution of the BAO dilation parameter is determined, we can evaluate our fiducial cosmology in Eq. (11) to convert our α_{iso} and α_{AP} constraints into distance measurements. When these constraints are passed onto the next stage of the pipeline to begin the inference of cosmological parameters, they are compressed into a Gaussian posterior that is only characterized by its mean and covariance matrix, extracted from the chains. Deviations from Gaussianity could impact the resulting cosmological constraints, so this assumption must be previously examined to ensure consistency.

Figure 9 shows the posterior distribution of BAO parameters from the DESI DR2 data. The mean of the marginalized posterior of each parameter has been

subtracted to ensure the constraints remain blinded. We overlay a multivariate Gaussian distribution with the same mean and covariance as the data. Overall, the data posterior shows excellent agreement with Gaussianity for all tracers. Some small deviations can be spotted at the 3σ contours, especially in tracers with lower S/N ratios such as the QSOs. However, we have explicitly verified that most of these small differences can be attributed to sampling noise from the Markov chain Monte Carlo chains. This figure also highlights the reduced parameter correlation in the $\alpha_{\text{iso}}\text{-}\alpha_{\text{AP}}$ plane compared with the constraints in the $\alpha_{\parallel}\text{-}\alpha_{\perp}$ basis.

VII. POST-UNBLINDING TESTS

While our fiducial choices were determined solely based on the validation process conducted with blinded data, we *also planned a set of post-unblinding tests in advance* to assess the stability of our results. These tests were designed before unblinding to prevent selection biases and ensure they did not influence the baseline analysis choices. These tests offer a complementary assessment of systematic uncertainties that may inform future refinements in the DESI analysis pipeline.

Below, we summarize the key post-unblinding tests conducted in this study.

A. Correlation in theoretical systematics

As in DESI DR1, our baseline analysis includes both statistical and systematic uncertainties, with systematic errors treated as uncorrelated across redshift bins and tracers. However, as a *sanity check*, we explore the impact of allowing these systematic uncertainties to be correlated across redshift bins and different tracers, testing whether such correlations significantly affect the inferred BAO constraints.

In both DR1 and DR2 analyses, we adopt a 0.1% theoretical systematic uncertainty on α_{iso} and a 0.2% uncertainty on α_{AP} , as shown in Table V. Since cross-tracer systematic correlations could arise from common theoretical assumptions—such as reconstruction modeling—we assess their impact by introducing correlations *only for theoretical systematics*, while keeping fiducial cosmology and HOD systematics uncorrelated. We evaluate two extreme cases:

- (1) *Fully correlated*: The theoretical systematic errors are assumed to be completely correlated across redshift bins and tracers, meaning that a systematic shift in one bin applies identically to all others. This represents the strongest correlation scenario we test.
- (2) *Partially correlated*: We introduce a moderate level of correlation by splitting the theoretical systematic uncertainty into half-correlated and half-uncorrelated components. Specifically, we assign a covariance of $1/2 \times (0.1\%)^2$ between all α_{iso} values across

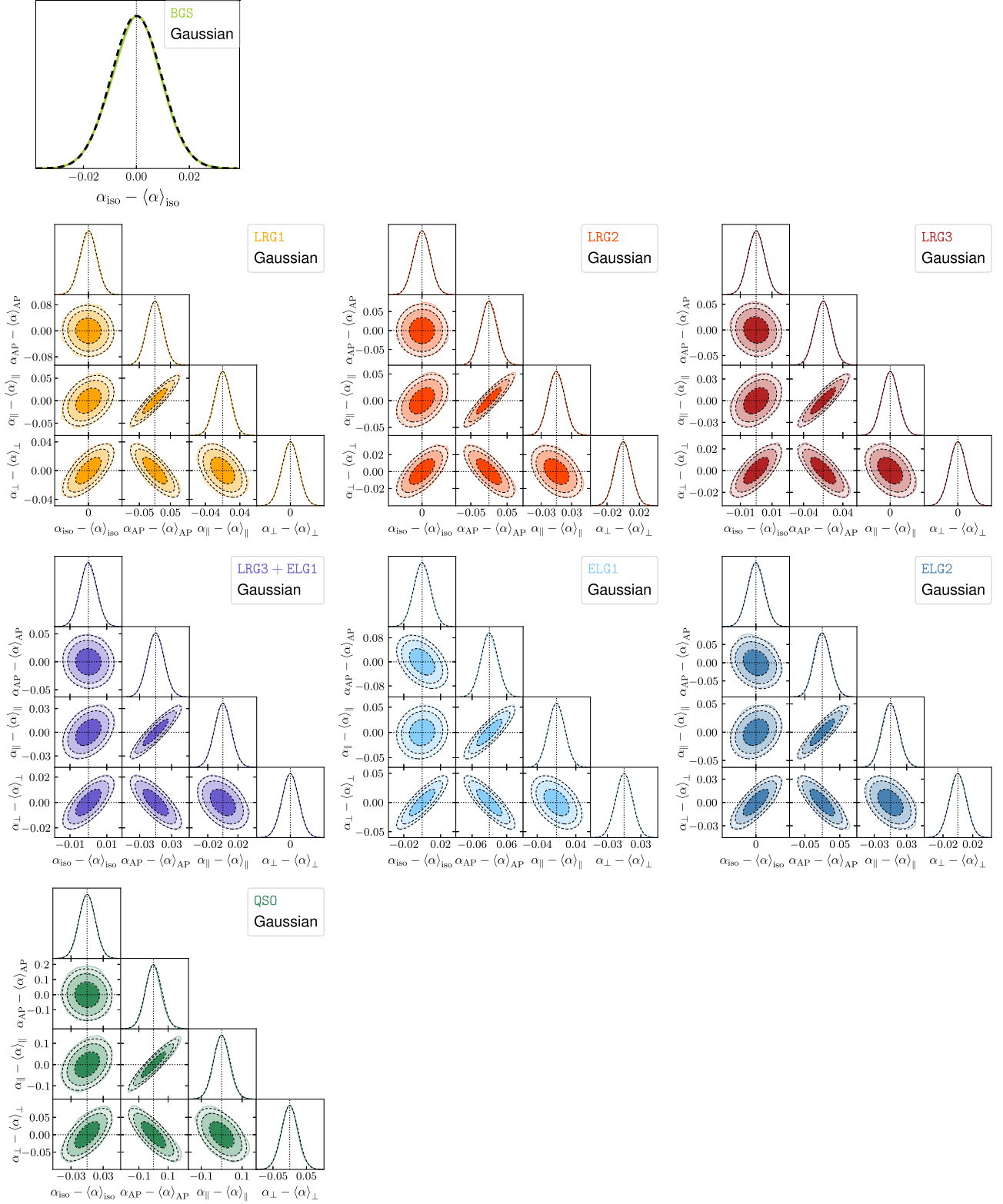


FIG. 9. Marginalized posterior distributions of BAO scaling parameters in two different bases, compared against a multivariate Gaussian with the same mean and covariance matrix. The inner, middle, and outer contours correspond to the 68%, 95%, and 99% confidence regions, respectively. The results demonstrate that the data posteriors exhibit excellent agreement with Gaussianity for all tracers.

redshift bins and tracers, and similarly, $1/2 \times (0.2\%)^2$ for α_{AP} values. This allows for some degree of coherence in the systematics while maintaining independent components.

By comparing these cases to our baseline, where systematic uncertainties remain uncorrelated, we assess whether correlated theoretical systematics could introduce biases in the BAO constraints. This treatment results in a *nondiagonal* systematics covariance matrix across different redshift bins, modifying the overall uncertainty structure when combined with statistical errors.

The effect of allowing correlated theoretical systematics on the inferred cosmology from the BAO distances is shown in Fig. 10. Note that the location of the data points and the size of the (diagonal) errors do not change in this figure. The solid black curve represents the standard w_0w_a CDM fit using DESI + CMB + DESY5, while the red dashed curve includes the impact of correlated

systematic uncertainties. We observe a negligible impact on the best fit w_0w_a CDM cosmology, which reflects our statistical errors dominating the systematic errors. The analysis of how correlated systematics are propagated into the cosmological parameter constraints is presented in Appendix B of the companion DESI cosmology analysis [37].

B. Galaxy bias in reconstruction

The blinded analysis limits us from determining an accurate bias for each tracer before unblinding. For DR2, we used the same baseline b_1 values as in DR1 for reconstruction, which were validated against survey validation (SV) [78–80,112] data and shown to be accurate within a 10% margin. For ELGs, we perform reconstruction across the entire redshift range (0.8–1.6) with a bias of 1.2. Due to the wide redshift range, the bias of the ELGs evolves slightly, such that the bias for ELG2 is close to or above 1.4 [30]. This represents more than a 10% offset from the baseline $b_1 = 1.2$ that we adopted for reconstructing ELGs, i.e., beyond the 10% tolerance we consider acceptable for optimal reconstruction.

After unblinding, therefore, we tested ELG2 with $b_1 = 1.4$ in reconstruction; we found only a 0.04% shift in α_{iso} and a 0.3% shift in α_{AP} with the same precisions as our baseline with $b_1 = 1.2$. Note that our pre-unblinding test with the ELG2 mocks, which had a clustering amplitude similar to the observed ELG2 (and therefore a bias similar to that of the data) was also analyzed with $b_1 = 1.2$ during reconstruction and we demonstrated that the results were unbiased (Table VII). These tests again prove the robustness of the BAO reconstruction against the assumption of galaxy bias.

C. BAO fits with different fiducial cosmologies

As described in Sec. IV, the fiducial cosmology assumption enters the pipeline in three stages: the redshift-to-distance conversion, leading to potential geometrical distortions; the sound horizon scale, due to fixing the linear power spectrum template; and the linear bias b_1 and growth rate f values assumed when running reconstruction.

In order to test the impact of the assumption of a fiducial cosmology, we repeated our baseline analysis with four additional sets of cosmological parameters. This test is analogous to the one presented in [89], where the fiducial-cosmology systematic error budget was quantified for the DR1 BAO results [27].

The tested cosmologies correspond to four alternative models from the AbacusSummit suite [75], each designed to probe different aspects of the fiducial cosmology assumption. These include a low Ω_m cosmology (c001), a thawing dark energy model with evolving $w(z)$ (c002), a scenario with a higher effective number of relativistic species, N_{eff} (c003), and a variation of our baseline cosmology with a lower matter fluctuation

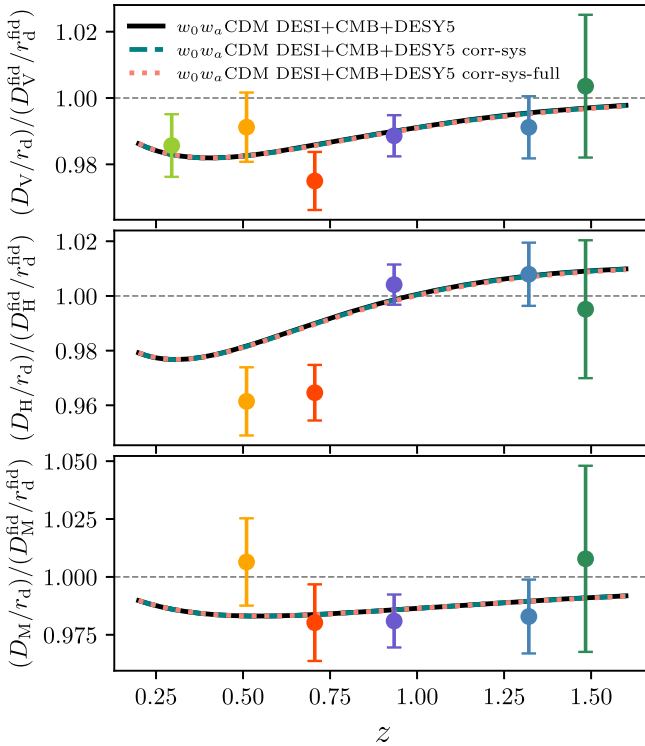


FIG. 10. Impact of correlated theoretical systematics on the cosmological inference from BAO distance measurements. Each panel shows the ratio of measured BAO distances relative to the fiducial cosmology, including D_V/r_d (top), D_H/r_d (middle), and D_M/r_d (bottom). The solid black curve represents the best-fit w_0w_a CDM model using DESI+CMB+DESY5. The teal dashed curve includes the additional effect of partially correlated systematic uncertainties, while the salmon dotted curve shows the scenario where theoretical systematics are fully correlated across redshift bins. The comparison between these cases quantifies the potential bias introduced by different assumptions about systematic correlations. Error bars include the systematic error budget in each redshift bin.

TABLE X. Secondary AbacusSummit cosmologies used to test the impact of the fiducial cosmology. Each column shows the ratio of a key cosmological quantity relative to our baseline fiducial cosmology. The third column quantifies the effect of the fiducial sound horizon ratio on the template, while the fourth and fifth columns display the range of expected geometrical distortions across the redshift interval $0.1 < z < 2.1$, with values varying as a function of redshift. Specifically, D_V and D_H/D_M are computed using the fiducial cosmology in each case, and their ratios with respect to the baseline cosmology are reported as ranges spanning the redshift interval. The last two columns indicate how the values of the growth rate f and the linear galaxy bias b_1 —both required for BAO reconstruction—change for each cosmology (again, showing their range across all tracer samples). The largest deviations in fiducial distances occur for the thawing dark energy case, where an isotropic rescaling of up to $\sim 7\%$ and an anisotropic distortion of up to $\sim 3\%$ are observed. In contrast, the low σ_8 cosmology maintains the same fiducial distances as the baseline cosmology, affecting only the galaxy bias assumptions for reconstruction.

Cosmology	Description	$r_d/r_d^{\text{baseline}}$	$D_V/D_V^{\text{baseline}}$	$\frac{D_H}{D_M} / \frac{D_H^{\text{baseline}}}{D_M^{\text{baseline}}}$	f/f^{baseline}	$b_1/b_1^{\text{baseline}}$
c001	$\Omega_m = 0.276$ Λ CDM	1.056	1.002–1.044	1.001–1.025	0.95–0.99	1.05–1.13
c002	$w_0 = -0.7$, $w_a = -0.5$ (thawing DE)	0.932	0.931–0.989	0.9713–0.995	0.98–1.00	0.95–1.07
c003	$N_{\text{eff}} = 3.70$ Λ CDM	1.022	1.001–1.019	1.001–1.011	0.98–1.00	0.95–1.07
c004	$\sigma_8 = 0.75$ Λ CDM	1.000	1.000–1.000	1.000–1.000	1.00	1.10–1.20

amplitude, σ_8 (c004). Notably, the BAO analysis has been shown to be robust beyond these models, for instance, in the context of modified gravity [113] and early dark energy models [114].

Each of these cosmologies impacts the BAO analysis in distinct ways, introducing different AP distortions across redshift, and modifying the sound horizon scale at the drag epoch due to shifts in the early-Universe expansion history. These effects are summarized in Table X, which provides a quantitative comparison of key cosmological parameters relative to the baseline.

In particular, the thawing-dark-energy cosmology introduces large geometrical distortions, with an expected isotropic contraction of distances of up to $\sim 7\%$ and anisotropic distortions of up to $\sim 3\%$ with respect to the baseline fiducial cosmology.

The high- N_{eff} cosmology has shown to cause shifts of up to $\sim 0.2\%$ in α_{iso} [89] due to the effect that changing the number of relativistic species has on the BAO phase shift in the linear power spectrum template. A similar shift had been observed before by [115], who tested differences of up to $\Delta N_{\text{eff}} = 2$ between synthetic data and the fitting template. As this effect stems from the assumed rescaling in the AP parameters with the sound horizon ratio [Eq. (11)], we regard it as a matter of interpretation. Nonetheless, if the cosmological analysis had indicated significant deviations from the fiducial N_{eff} value, we would have performed a reanalysis using an updated fiducial cosmology, as planned before unblinding. However, since no such deviations were observed, this was not required.

The low- σ_8 case, on the other hand, primarily influences the bias used in BAO reconstruction, as this cosmology is otherwise identical to our baseline cosmology except for a low A_s (and consequently a low σ_8) value.

For a consistent analysis across different fiducial models, we updated the reconstruction pipeline by adjusting the values of the linear bias (b_1) and growth rate (f)

accordingly. Additionally, we recomputed the RascalC covariance matrices to reflect the altered redshift-to-distance conversion and reconstruction parameters for each tested cosmology.

Figure 11 shows the measurements obtained for D_V/r_d and D_H/D_M in comparison to the values reported in [37]. Across all samples, the deviations from the baseline measurements are relatively small, with the largest differences observed for BGS. In all cases, the variations are consistent with zero, staying within 0.5σ . This confirms the stability of the BAO fits against the choice of fiducial cosmology.¹⁹

Pipeline reanalysis for w_0w_a CDM cosmology: We repeated the entire analysis pipeline using the best-fit w_0w_a CDM cosmology from DESI + CMB + DESY5 data, as reported in Sec. VII of [37], to assess the impact of the fiducial cosmology choice on the final BAO measurements. The results of this reanalysis, shown in Fig. 11, demonstrate that the inferred BAO distances remain highly consistent with those obtained using the baseline Λ CDM fiducial cosmology. The largest deviation in D_H/D_M occurs for LRG1 with 0.4σ , while for D_V/r_d , the largest relative difference is 0.2σ , observed in the ELG2 bin.

D. DESI-SDSS BAO comparison at $z \sim 0.7$

We compare the BAO measurements from DESI DR1, DESI DR2, and SDSS at $z \sim 0.7$,²⁰ where the DR1 analysis reported a 2.5σ – 3σ tension with SDSS. To ensure a

¹⁹We emphasize that the tests presented in this section serve as a null test on the real data; they assess whether the recovered BAO parameters change when analyzed with different fiducial cosmologies. In contrast, the systematic error budget discussed in Sec. V B is based on mock catalogs and quantifies the expected bias from such variations after marginalizing over sample variance.

²⁰We consider the slight difference in effective redshift ($\Delta z = 0.008$) between DESI and SDSS bins by multiplying $D^{\text{fid}}(z_{\text{DESI}})/D^{\text{fid}}(z_{\text{SDSS}})$, the ratio of distance in the fiducial cosmology, by the SDSS BAO measurement.

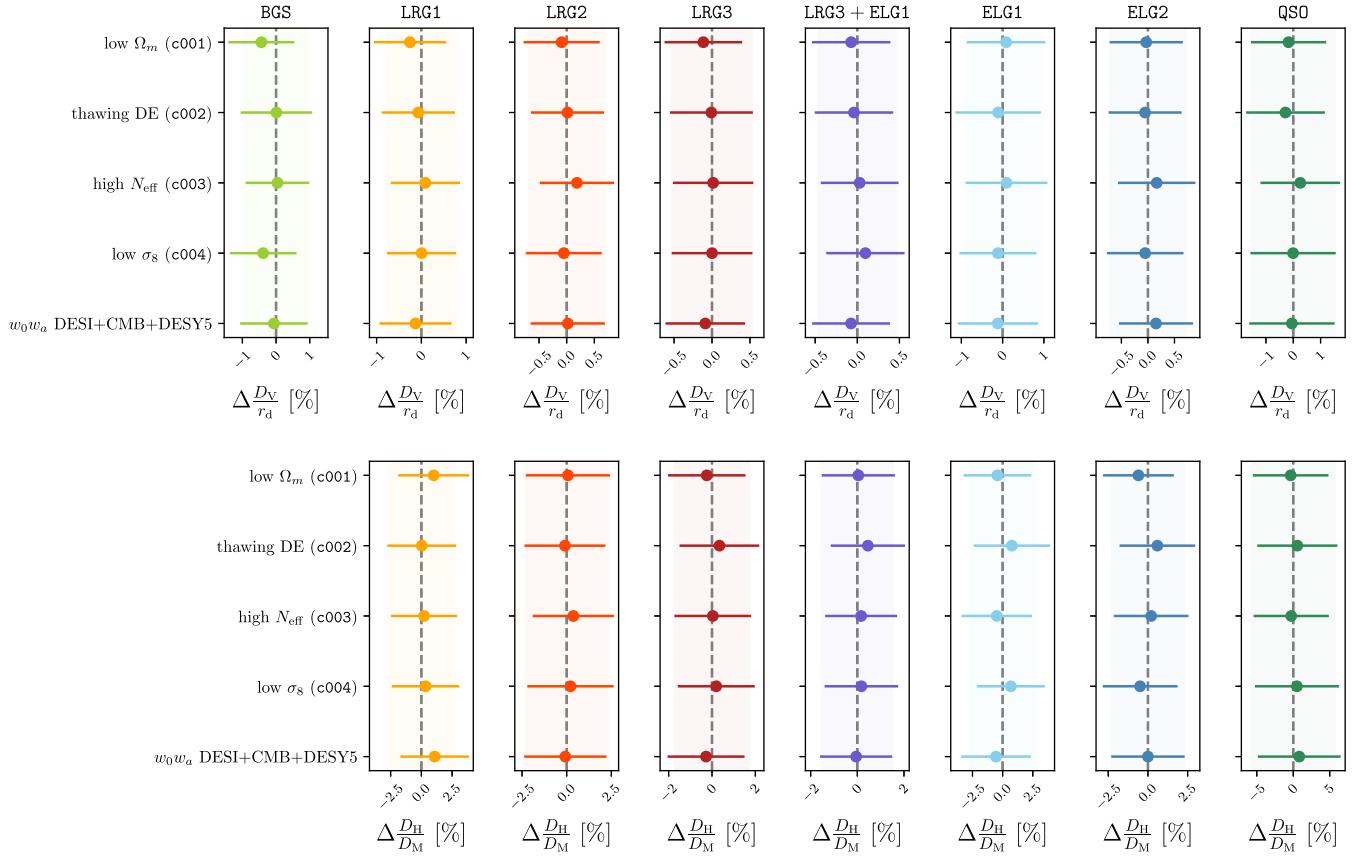


FIG. 11. Comparison of the distance measurements obtained by repeating the analysis with different fiducial cosmologies. The differences are taken with respect to the values reported in [37]. The error bars correspond to the uncertainties in the distances measured for each fiducial cosmology, whereas the shaded regions represent the errors from the baseline analysis.

consistent comparison, we use the DESI reanalysis of SDSS, which applies the DESI pipeline—including reconstruction, clustering measurements, and BAO fitting to the eBOSS LRG catalog. This reanalysis, presented in Table 17 of [27], yields very similar results to the published eBOSS measurements but removes pipeline-induced differences.

As shown in Fig. 12, the DR2 measurement lies between DR1 and SDSS, consistent with the interpretation from the DR1 paper that the observed tension is likely due to statistical fluctuations. The improved DR2 precision further supports this conclusion, as it reduces the statistical scatter while remaining broadly consistent with both data sets. A more detailed quantitative comparison can be found in Sec. III C 2 of [37].

VIII. SUMMARY AND CONCLUSIONS

We have presented the validation of the DESI DR2 measurements of baryon acoustic oscillations from galaxies and quasars. Together with the BAO measurements from the Ly α forest [116], these constitute the main data set that was used for cosmological inference in [37]. The validation process involved extensive testing with mock galaxy catalogs that mimic the selection and clustering properties

of different galaxy tracers, as well as analyses using a blinded version of the DESI DR2 LSS catalogs, followed by complementary post-unblinding checks.

Our analysis pipeline remains largely the same as that used for the DESI DR1 BAO analysis [27], with minor modifications to account for the expanded data set and refinements in analysis choices. These include adjustments to the BGS magnitude cut, an increased minimum fitting scale for BAO measurements, the adoption of 2D BAO fits for select tracers and additional data-splitting tests to assess systematic effects (detailed in Sec. V).

We have also outlined a series of rigorous tests to determine whether we were ready to finalize our pipeline settings and unblind the data for the official BAO cosmology analysis. These tests provided key validation steps, ensuring that our results were robust against potential sources of systematic bias (detailed in Sec. VIB). This unblinding checklist, detailed in Table VI, was successfully completed, leading to the unblinding of the DR2 BAO data during the DESI collaboration meeting in December 2024.

Our main findings can be summarized as follows:

- (1) We find strong agreement in clustering amplitude between DR1 and DR2 samples, with DR2 benefiting from a significant reduction in statistical scatter

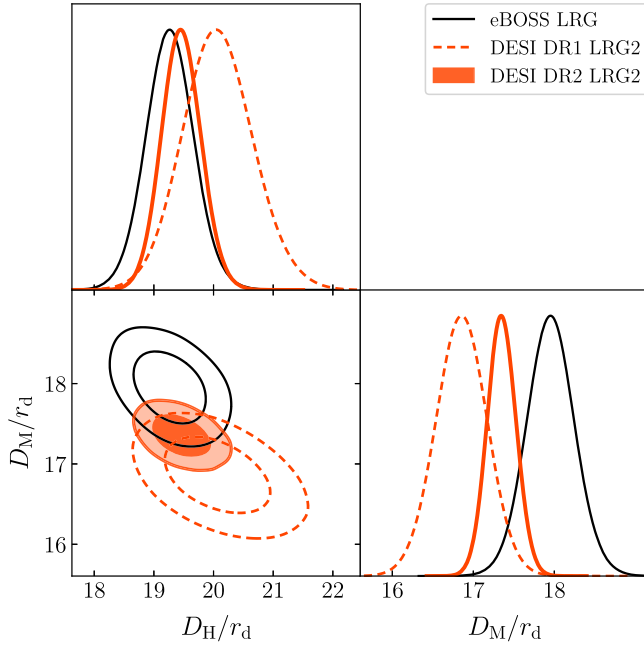


FIG. 12. Comparison of BAO constraints from DESI DR1, DESI DR2, and SDSS at $z \sim 0.7$. The plot shows the 68% and 95% confidence contours for the LRG2 bin in the D_H/r_d - D_M/r_d plane. The black solid contours represent the eBOSS (SDSS) measurement, the orange dashed contours show the DESI DR1 result, and the filled orange contours correspond to DESI DR2. The DR2 result lies between DR1 and SDSS, supporting the interpretation that the previously observed tension between DR1 and SDSS is likely a statistical fluctuation.

due to its larger volume. The improvement in precision is particularly pronounced for the BGS and ELG samples, as shown in Fig. 4. The combined BAO precision, including both statistical and systematic uncertainties, has improved from $\sim 0.52\%$ in DR1 to 0.301% in DR2, corresponding to a factor of 1.73 improvement. The impact of systematics remains modest, with $\sigma_{\text{stat+syst}} < 1.06\sigma_{\text{stat}}$, indicating that the results are still firmly statistics-dominated. Furthermore, DESI DR2 achieves a factor of 2.18 improvement over previous SDSS BAO measurements, which had a precision of $\sim 0.62\%$, highlighting the substantial gain in precision enabled by DESI's expanded data set.

- (2) The DESI DR2 BAO constraints are robust across multiple variations in analysis choices, including different data vectors (correlation function vs power spectrum, 1D vs 2D BAO fit), modeling (different broadband parametrizations and parameter priors), and treatment of observational systematics. We tested the response of the constraints across multiple data splits, including redshift bins, sky regions, and magnitude cuts, consistently finding agreement with the baseline (see Fig. 8).

- (3) Density-field reconstruction significantly enhances the precision of BAO constraints for all DESI tracers (Fig. 5). The uncertainty on the isotropic (anisotropic) BAO scaling parameter α_{iso} (α_{AP}) is reduced by up to 42% (47%) compared with results using unreconstructed catalogs. Compared with DR1, reconstruction efficiency has improved for almost all tracers, likely due to smaller boundary effects and higher survey completeness in DR2, which allow for better displacement field estimation. This level of improvement aligns with expectations from mock catalog analyses and further demonstrates the effectiveness of reconstruction in reducing noise and enhancing the BAO signal.
- (4) The posterior distribution of BAO scaling parameters is well approximated by a Gaussian for all tracers (Fig. 9), enabling the posterior to be fully characterized by a mean and covariance matrix for subsequent cosmological inference.
- (5) BAO constraints remain stable under different fiducial cosmologies. We tested the impact of our fiducial cosmology assumptions by reanalyzing BAO measurements using different cosmologies, including a low- Ω_m model, a thawing dark energy model, a high- N_{eff} scenario, and a low- σ_8 cosmology (Fig. 11). Across all tracers, deviations in BAO constraints remain within 0.4σ , confirming that our measurements are not significantly biased by fiducial cosmology choices. The largest variations appear in the BGS sample, but even these remain well within statistical uncertainties.
- (6) DESI DR2 provides a significantly improved BAO measurement compared with SDSS. We examined the reported tension between DESI DR1 and SDSS BAO measurements at $z \sim 0.7$, which was previously found to be at the 2.5σ – 3σ level. Using the DESI pipeline to reanalyze the eBOSS LRG sample, we found that the reanalysis closely matches the published SDSS results, confirming that differences in pipeline methodology were not responsible for the tension. As shown in Fig. 12, the DESI DR2 result lies between DR1 and SDSS, supporting the interpretation from the DR1 analysis that the tension was likely due to statistical fluctuations rather than a systematic discrepancy.
- (7) Correlated systematics do not significantly impact the BAO results. As a post-unblinding test, we investigated the potential impact of correlated systematics across redshift bins and tracers. We considered two extreme scenarios: one where theoretical systematics were fully correlated across bins, and another where they were only partially correlated (Fig. 10). In both cases, the deviations from the baseline analysis are minimal.

The validated DESI DR2 BAO data set represents a major step forward in precision cosmology. Compared with DR1, it provides more stringent constraints, forming a robust foundation for upcoming studies on dark energy, modified gravity, and the Universe’s expansion history. Future work will extend this analysis by refining systematics control and testing alternative modeling approaches to further enhance the precision and reliability of DESI’s BAO measurements. These results mark a crucial milestone for DESI’s mission, setting the stage for even more precise constraints with future data releases and strengthening the role of BAO as a cornerstone of modern cosmology.

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Any opinions, findings, and conclusions or recommendations expressed in this material are those of the author(s) and do not necessarily reflect the views of the U.S. National Science Foundation, the U.S. Department of Energy, or any of the listed funding agencies.

DATA AVAILABILITY

The data that support the findings of this article are openly available [118].

APPENDIX A: BAO MEASUREMENTS IN FS

Parallel to the analysis in CS, we also performed the analysis in FS. In this appendix, we highlight the main results obtained in FS compared with the ones we obtained in the fiducial CS analysis.

In Fig. 13, we show the isolated BAO feature in the multipole moments of the DESI DR2 power spectra of galaxies and quasars. The upper subpanels display the monopole, while the lower subpanels show the quadrupole. The best-fit model is shown as solid lines except for the quadrupole of the BGS tracer, since for this tracer the baseline fit is a 1D fit and we show it with a dashed line.

After measuring the power spectrum for the different tracers, we proceeded to run the BAO fits. The settings we used are displayed in Table IV.

In Fig. 14 we show the pairwise comparison of the CS and FS BAO-fit results for the unblinded data. In the panel on top, we show differences in α_{iso} , and in the one on the bottom, α_{AP} . We note that the baseline case is a 1D BAO fit for BGS and a 2D fit for the others (which is the reason why BGS is not shown in the panel on the bottom). The colored-shaded areas show the 1σ regions of the baseline case, and the gray-shaded ones show one-third of 1σ . We see that CS and FS results are in very good agreement: the dashed lines at $\Delta\alpha = 0$ lie within the gray-shaded areas for most tracers, and the different cases (1D fit, pre-reconstruction, NGC and SGC) are typically consistent with zero (SGC is usually the case that differs the most due to its lower constraining power). The largest differences between CS and FS are for the BGS tracer, but they are still consistent with zero at the 1σ level (besides the SGC case, which is slightly outside the 1σ area of the baseline case).

We also ran the BAO fits on the Abacus-2 DR2 mocks. This allowed us to check whether the differences between CS and FS found in the unblinded data (which we already showed in Fig. 14) were consistent with those of our simulations. In Fig. 15 we show the differences in α_{iso} between CS and FS for the different tracers, for both the mocks and unblinded data. We see that the results in the data (dashed black lines) are usually well within the distribution of the mocks (histograms). However, this is not the case for BGS, for which the unblinded data shows a 0.55% offset, which is a bit larger than the most extreme case from the mocks. An offset of such magnitude is also observed in other tracers with higher S/N, such as ELG1. The BGS mock distribution is somewhat asymmetric, with an extended tail on the left, making such a deviation plausible, especially given the limited number of mocks available for this test. In absolute value, two mocks have larger $\Delta\alpha_{\text{iso}}$ than the unblinded data, i.e., two mocks on the left-hand side of the tail of the distribution for BGS in Fig. 15 (first panel) have a more extreme value than the one found on the data (the dashed black line). Considering that the offset is within an acceptable range and does not exceed

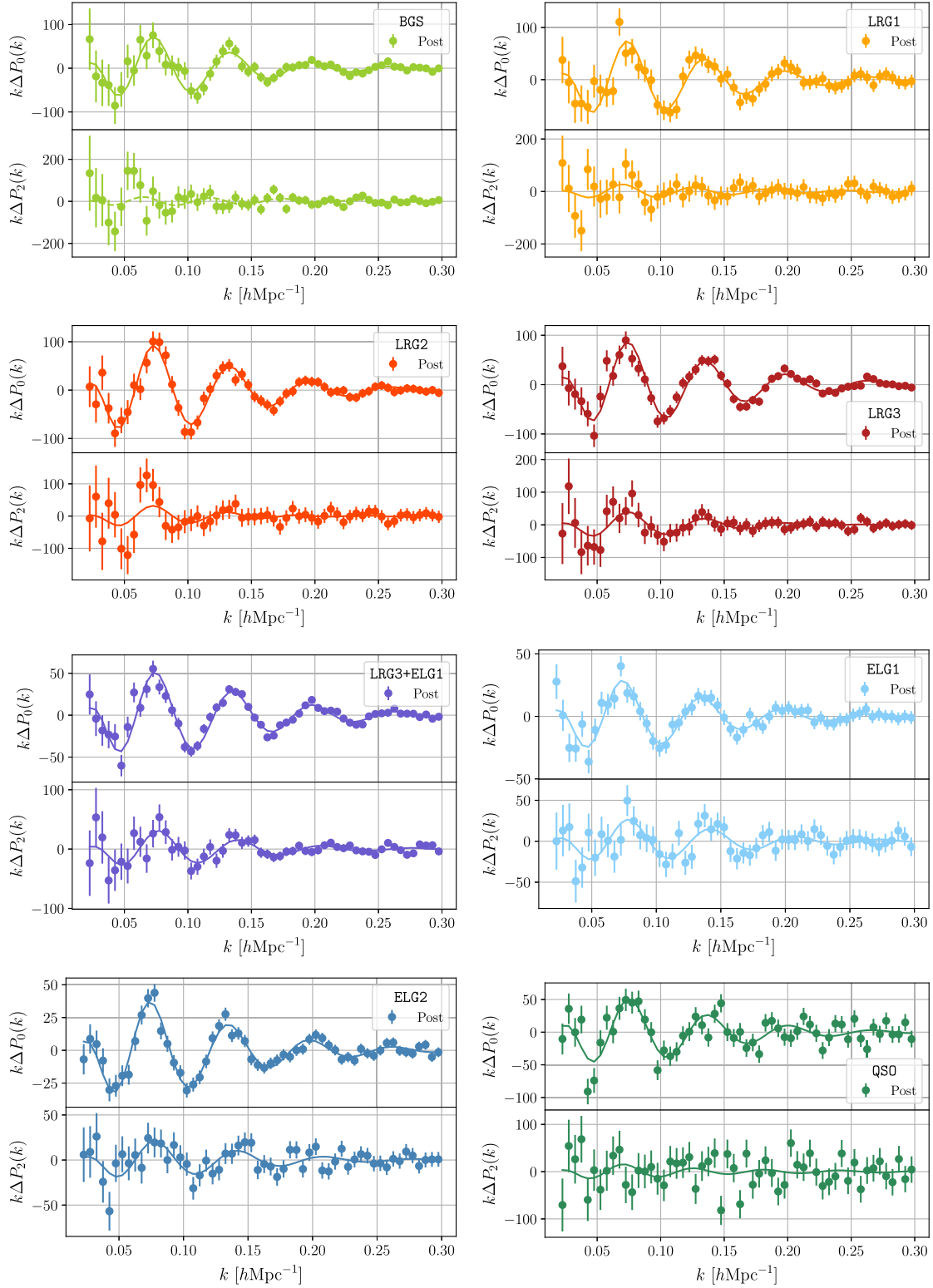


FIG. 13. The isolated BAO feature in the monopole and quadrupole moments of the DESI DR2 power spectrum for different tracers. The filled circles represent the post-reconstruction measurements, while the solid lines correspond to the best-fit BAO model. The upper and lower subpanels in each plot display the monopole and quadrupole moments, respectively. For BGS, the baseline fit is a 1D fit and we show the best-fit model for the quadrupole as a dashed line instead of a solid one. All other tracers use a 2D fit that incorporates both the monopole and quadrupole moments as the baseline settings. Error bars denote 68% confidence intervals. The y axis units of $h^{-3} \text{ Mpc}^3$ are omitted for clarity.

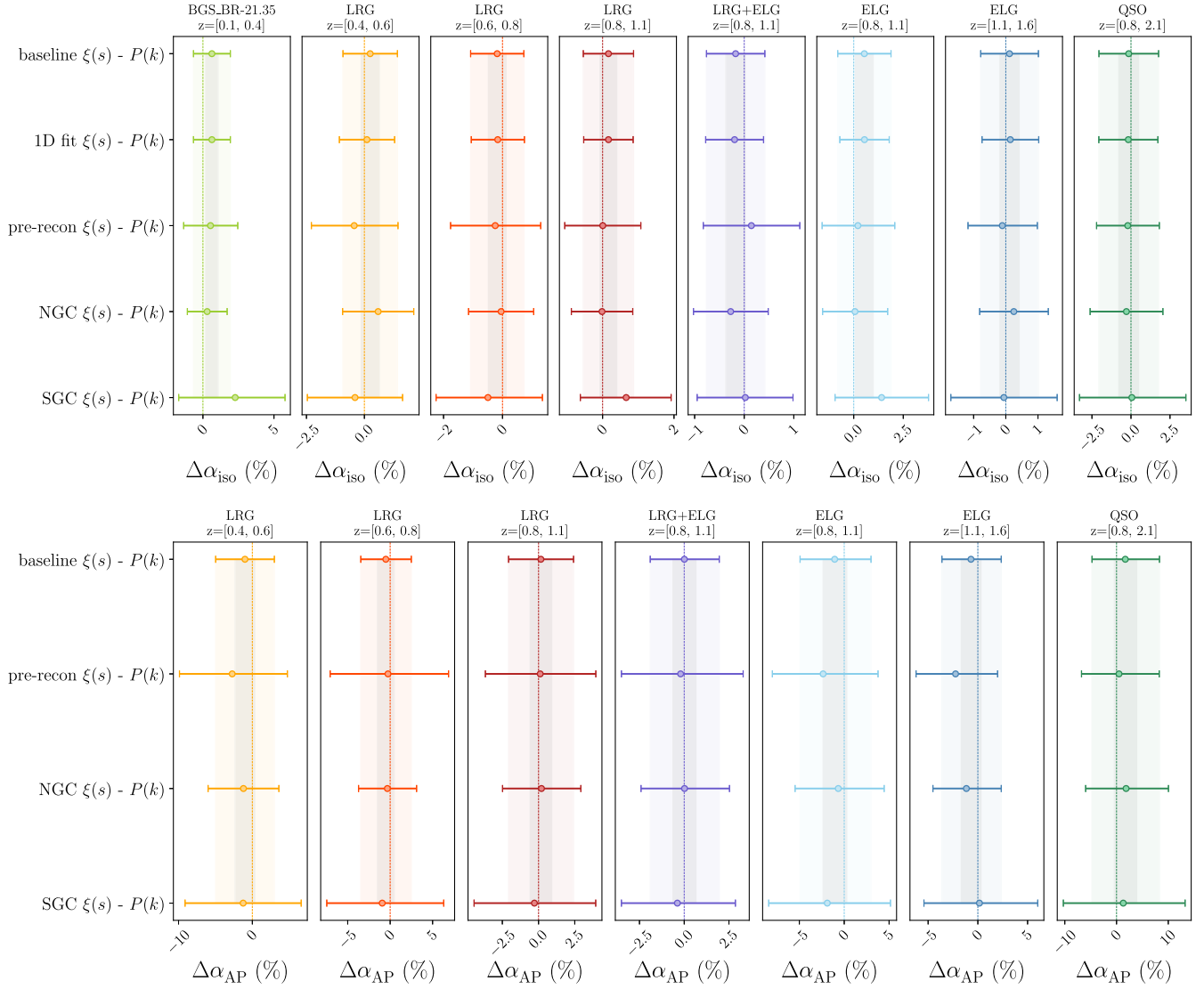


FIG. 14. Pairwise comparison of the correlation function and the power spectrum BAO fit results for the unblinded data. In the panel on top, we show differences in α_{iso} , and in the one on the bottom, α_{AP} . We note that the baseline case is a 1D BAO fit for BGS and a 2D fit for the others (which is the reason why BGS is not shown in the panel on the bottom). The colored-shaded areas show the 1σ regions of the baseline case, and the gray-shaded ones show one-third of 1σ .

3σ , we conclude that the BGS baseline CS BAO fit remains valid.

APPENDIX B: ASSESSING 1D VS 2D BAO FITS FOR BGS, ELG, AND QSO

BAO measurements can be performed using either 1D or 2D fits. A 1D fit constrains only the isotropic BAO scaling parameter, α_{iso} , while a 2D fit simultaneously constrains both the isotropic and anisotropic scaling parameters, α_{iso} and α_{AP} . While 2D fits provide additional cosmological information by capturing anisotropic distortions, they

require sufficiently high S/N and stable posterior distributions to ensure robust results.

In DESI DR1, the BAO fits for BGS, ELG1, and QSOs were performed using 1D fits only, as the statistical precision of the data was not sufficient to justify a 2D fit. With DESI DR2, the increased survey volume and improved statistical precision have enhanced the feasibility of 2D BAO fits for these tracers. Initial mock-based tests suggested that the DR2 data set has sufficient S/N to perform 2D BAO fits for these samples. However, a detailed examination of the fits was required to ensure their robustness. In DESI DR2, the choice between 1D and

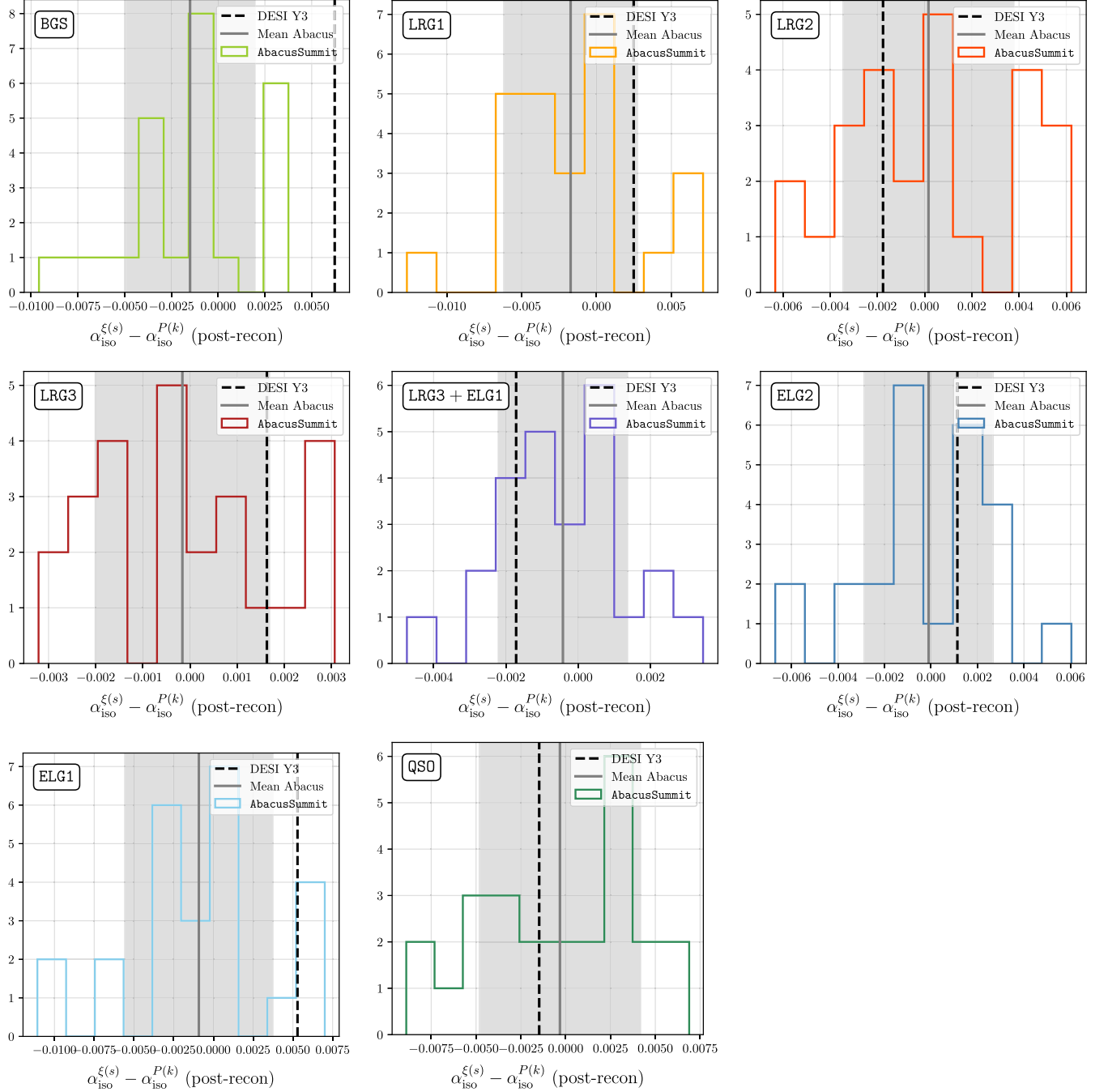


FIG. 15. The different panels show the distribution of $\Delta\alpha_{\text{iso}}$ between CS and FS in the mocks (histograms) and the results obtained on the unblinded DESI DR2 data (dashed black lines). The shaded gray areas show the 1σ region around the mean of the mocks, and the center is highlighted with a gray line. The first panel shows the results of the 1D BAO fit on the BGS tracer, whereas we show the results of the 2D fit for all the other tracers.

2D fits was evaluated using a combination of mock-based tests and validation metrics on the blinded data.

1. Criteria for selecting 1D vs 2D BAO fits

The choice between 1D and 2D fits was guided by the following considerations:

- (1) The statistical precision of the sample: A higher S/N allows for stable 2D fits, while a low S/N may introduce bias and unstable parameter estimates.
- (2) The Gaussianity of the posterior distributions, particularly for the anisotropic parameter α_{AP} , avoiding complications of including a weak non-Gaussian constraint on α_{AP} .

- (3) The correlation between α_{iso} and α_{AP} in the 2D fit should match mock expectations and align with theoretical predictions for optimal BAO isolation (Seo and Eisenstein [8]).
- (4) The distribution of χ^2 values when compared with mock realizations.

2. Evaluation of 2D fits in DR2

Tests performed on the Abacus-2 DR2 mocks suggested that we have sufficient S/N to apply 2D BAO fits to BGS, ELG1, and QSOs. The improvements in DR2 relative to DR1 suggested that 2D fits could be feasible for these tracers. However, a closer examination of the results led to a more nuanced decision:

a. BGS

For BGS, the 2D fit is borderline acceptable—it formally passes validation tests but raises concerns:

- (1) Non-Gaussianity in α_{AP} : The posterior distribution for α_{AP} exhibits elongation and deviations from Gaussianity, suggesting a more complex likelihood structure that may not be well approximated by a simple Gaussian assumption.
- (2) High $\alpha_{\text{iso}} - \alpha_{\text{AP}}$ correlation: The blinded data shows $r = 0.48$, unusually high compared with other tracers. In the $\alpha_{\parallel} - \alpha_{\perp}$ basis, it lies slightly outside the expected mock distribution ($r = -0.49$ vs -0.42 to -0.47). Theoretical predictions [119] suggest an optimal BAO isolation should yield a correlation

value of $r \sim -0.41$, which holds for other tracers but not for BGS.

At low redshifts ($z < 0.4$), 2D BAO fits are inherently challenging due to:

- (1) Weak AP distortions: The AP effect—driven by D_H/D_M —is weaker at low z , reducing sensitivity to cosmology.
- (2) Lower S/N in anisotropic clustering: RSDs remain strong, but distinguishing AP effects from RSDs and statistical noise is harder.
- (3) Limited number density: The lower BGS number density further reduces the precision of α_{AP} measurements.

Given these factors, the 1D BAO fit—which measures only α_{iso} —is a more stable and robust choice for BGS. While a higher S/N data set could, in principle, constrain the AP effect, the limited statistical power and elevated $\alpha_{\text{iso}} - \alpha_{\text{AP}}$ correlation make the 2D fit less reliable.

Although the BGS 2D fit is reasonable and passes most consistency tests, minor deviations and potential non-Gaussian likelihood effects led us to conservatively adopt the 1D fit as our default. Importantly, both the 1D and 2D fits yield consistent results, reinforcing the robustness of the BAO measurement for BGS.

b. ELG1

The 2D fits for ELG1 show hints of similar behavior as BGS, particularly in the increased correlation between α_{iso} and α_{AP} . However, the differences are less pronounced, and the fit remains within the expected distribution of mocks.

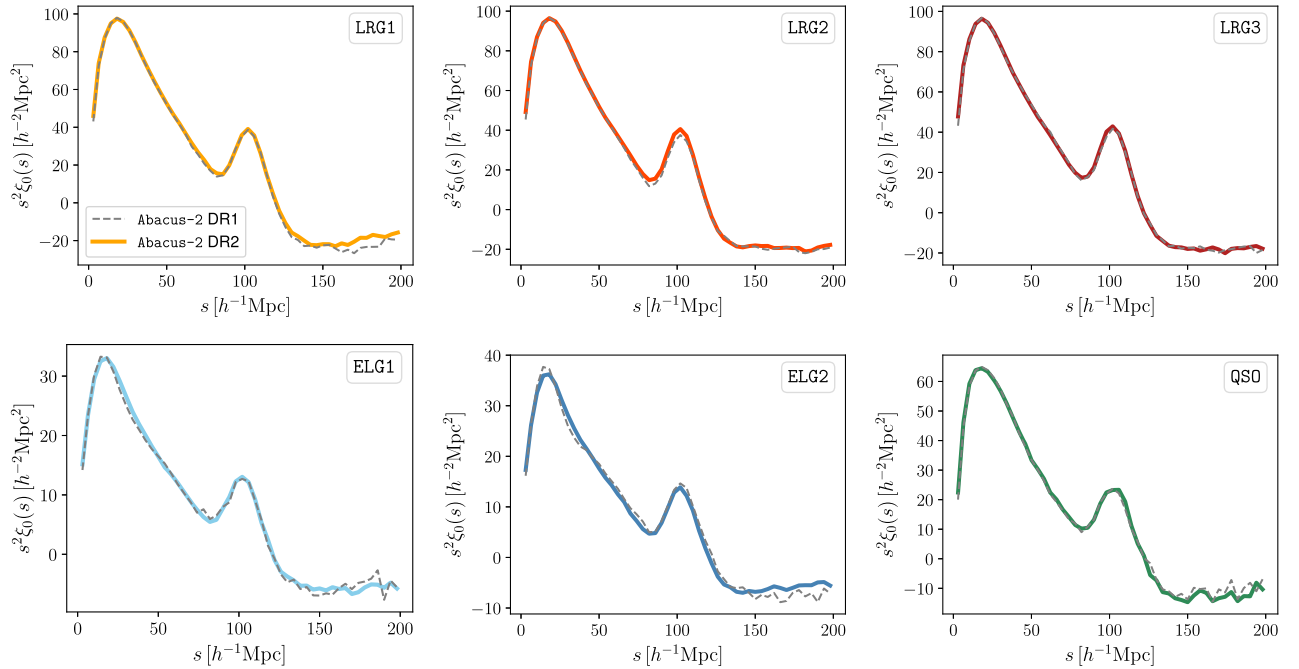


FIG. 16. Comparison of the monopole of the two-point correlation function, $\xi_0(s)$, between the Abacus-2 DESI DR1 mocks (gray) and the Abacus-2 DESI DR2 mocks (colored) for various tracer samples and redshift ranges (we show the average over the 25 mocks). The panels show results for LRG ($0.4 < z < 0.6$, $0.6 < z < 0.8$, $0.8 < z < 1.1$), ELG ($0.8 < z < 1.1$, $1.1 < z < 1.6$), and QSOs ($0.8 < z < 2.1$).

Since ELG1 is not used for cosmological inference (we use the combined LRG3 + ELG1 sample instead), we do not need to make a strict decision. We provide both 1D and 2D fits for ELG1, but only use the LRG3 + ELG1 combination for final cosmological constraints.

c. QSO

The 2D fits for QSOs pass all validation tests, including consistency with mocks, Gaussianity of posteriors, and robustness against data splits. The improvement in S/N at high redshift ensures that α_{AP} is well constrained, making a 2D fit the preferred choice. Therefore, we use the 2D fit as the default for QSOs.

APPENDIX C: CONSISTENCY BETWEEN DESI DR1 AND DR2 ABACUS-2 MOCKS

In Fig. 4 we show the measurements of the two-point correlation functions of the DESI DR2 compared with those of the DESI DR1. In this appendix, we show analogous results but for the Abacus-2 mocks. In Fig. 16 we show the monopole of the two-point correlation function measured in the Abacus-2 mocks (averaged over the 25 mocks) for the different tracers and redshift bins. Solid-colored lines show the measurements on the Abacus-2 DR2 mocks, whereas the shaded-gray ones show DR1. We see that the results are consistent between the two Abacus-2 data sets.

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- [1] D. J. Eisenstein and W. Hu, *Astrophys. J.* **496**, 605 (1998).
 - [2] D. J. Eisenstein, *New Astron. Rev.* **49**, 360 (2005).
 - [3] B. Bassett and R. Hlozek, in *Dark Energy: Observational and Theoretical Approaches*, edited by P. Ruiz-Lapuente (Cambridge University Press, Cambridge, 2010), p. 246.
 - [4] D. H. Weinberg, M. J. Mortonson, D. J. Eisenstein, C. Hirata, A. G. Riess, and E. Rozo, *Phys. Rep.* **530**, 87 (2013).
 - [5] D. J. Eisenstein, I. Zehavi, D. W. Hogg, R. Scoccimarro, M. R. Blanton, R. C. Nichol, R. Scranton, H.-J. Seo, M. Tegmark, Z. Zheng, S. F. Anderson, J. Annis, N. Bahcall, J. Brinkmann, S. Burles *et al.*, *Astrophys. J.* **633**, 560 (2005).
 - [6] W. J. Percival, C. M. Baugh, J. Bland-Hawthorn, T. Bridges, R. Cannon, S. Cole, M. Colless, C. Collins, W. Couch, G. Dalton, R. De Propris, S. P. Driver, G. Efstathiou, R. S. Ellis, C. S. Frenk *et al.*, *Mon. Not. R. Astron. Soc.* **327**, 1297 (2001).
 - [7] S. Cole, W. J. Percival, J. A. Peacock, P. Norberg, C. M. Baugh, C. S. Frenk, I. Baldry, J. Bland-Hawthorn, T. Bridges, R. Cannon, M. Colless, C. Collins, W. Couch, N. J. G. Cross, G. Dalton *et al.*, *Mon. Not. R. Astron. Soc.* **362**, 505 (2005).
 - [8] H.-J. Seo and D. J. Eisenstein, *Astrophys. J.* **633**, 575 (2005).
 - [9] D. J. Eisenstein, H.-J. Seo, E. Sirko, and D. N. Spergel, *Astrophys. J.* **664**, 675 (2007).
 - [10] N. Padmanabhan, X. Xu, D. J. Eisenstein, R. Scalzo, A. J. Cuesta, K. T. Mehta, and E. Kazin, *Mon. Not. R. Astron. Soc.* **427**, 2132 (2012).
 - [11] C. Blake, T. Davis, G. B. Poole, D. Parkinson, S. Brough, M. Colless, C. Contreras, W. Couch, S. Croom, M. J. Drinkwater, K. Forster, D. Gilbank, M. Gladders, K. Glazebrook, B. Jelliffe *et al.*, *Mon. Not. R. Astron. Soc.* **415**, 2892 (2011).
 - [12] C. Blake, E. A. Kazin, F. Beutler, T. M. Davis, D. Parkinson, S. Brough, M. Colless, C. Contreras, W. Couch, S. Croom, D. Croton, M. J. Drinkwater, K. Forster, D. Gilbank, M. Gladders *et al.*, *Mon. Not. R. Astron. Soc.* **418**, 1707 (2011).
 - [13] E. A. Kazin, J. Koda, C. Blake, N. Padmanabhan, S. Brough, M. Colless, C. Contreras, W. Couch, S. Croom, D. J. Croton, T. M. Davis, M. J. Drinkwater, K. Forster, D. Gilbank, M. Gladders *et al.*, *Mon. Not. R. Astron. Soc.* **441**, 3524 (2014).
 - [14] L. Anderson, E. Aubourg, S. Bailey, D. Bizyaev, M. Blanton, A. S. Bolton, J. Brinkmann, J. R. Brownstein, A. Burden, A. J. Cuesta, L. A. N. da Costa, K. S. Dawson, R. de Putter, D. J. Eisenstein, J. E. Gunn *et al.*, *Mon. Not. R. Astron. Soc.* **427**, 3435 (2012).
 - [15] L. Anderson, É. Aubourg, S. Bailey, F. Beutler, V. Bhardwaj, M. Blanton, A. S. Bolton, J. Brinkmann, J. R. Brownstein, A. Burden, C.-H. Chuang, A. J. Cuesta, K. S. Dawson, D. J. Eisenstein, S. Escoffier *et al.*, *Mon. Not. R. Astron. Soc.* **441**, 24 (2014).
 - [16] S. Alam, M. Ata, S. Bailey, F. Beutler, D. Bizyaev, J. A. Blazek, A. S. Bolton, J. R. Brownstein, A. Burden, C.-H. Chuang, J. Comparat, A. J. Cuesta, K. S. Dawson, D. J. Eisenstein, S. Escoffier *et al.*, *Mon. Not. R. Astron. Soc.* **470**, 2617 (2017).
 - [17] M. Ata, F. Baumgarten, J. Bautista, F. Beutler, D. Bizyaev, M. R. Blanton, J. A. Blazek, A. S. Bolton, J. Brinkmann, J. R. Brownstein, E. Burtin, C.-H. Chuang, J. Comparat, K. S. Dawson, A. de la Macorra *et al.*, *Mon. Not. R. Astron. Soc.* **473**, 4773 (2018).
 - [18] S. Alam, M. Aubert, S. Avila, C. Bland, J. E. Bautista, M. A. Bershad, D. Bizyaev, M. R. Blanton, A. S. Bolton, J. Bovy, J. Brinkmann, J. R. Brownstein, E. Burtin, S. Chabanier, M. J. Chapman *et al.*, *Phys. Rev. D* **103**, 083533 (2021).
 - [19] M. White, in *The Davis Meeting on Cosmic Inflation* (2003), p. 18, [arXiv:astro-ph/0305474](https://arxiv.org/abs/astro-ph/0305474).
 - [20] P. McDonald and D. J. Eisenstein, *Phys. Rev. D* **76**, 063009 (2007).
 - [21] T. M. C. Abbott, M. Adamow, M. Agüena, S. Allam, O. Alves, A. Amon, F. Andrade-Oliveira, J. Asorey, S. Avila, D. Bacon, K. Bechtol, G. M. Bernstein, E. Bertin, J. Blazek, S. Bocquet *et al.*, *Phys. Rev. D* **110**, 063515 (2024).

- [22] M. Levi, C. Bebek, T. Beers, R. Blum, R. Cahn, D. Eisenstein, B. Flaugher, K. Honscheid, R. Kron, O. Lahav, P. McDonald, N. Roe, and D. Schlegel (DESI Collaboration), [arXiv:1308.0847](#).
- [23] A. Aghamousa, J. Aguilar, S. Ahlen, S. Alam, L. E. Allen, C. Allende Prieto, J. Annis, S. Bailey, C. Balland, O. Ballester, C. Baltay, L. Beaufore, C. Bebek, T. C. Beers *et al.* (DESI Collaboration), [arXiv:1611.00036](#).
- [24] A. Dey, D. J. Schlegel, D. Lang, R. Blum, K. Burleigh, X. Fan, J. R. Findlay, D. Finkbeiner, D. Herrera, S. Juneau, M. Landriau, M. Levi, I. McGreer, A. Meisner, A. D. Myers *et al.*, *Astron. J.* **157**, 168 (2019).
- [25] A. G. Adame, J. Aguilar, S. Ahlen, S. Alam, G. Aldering, D. M. Alexander, R. Alfarsy, C. Allende Prieto, M. Alvarez, O. Alves, A. Anand, F. Andrade-Oliveira, E. Armengaud, J. Asorey *et al.* (DESI Collaboration), *Astron. J.* **168**, 58 (2024).
- [26] M. Abdul-Karim, A. G. Adame, D. Aguado, J. Aguilar, S. Ahlen, S. Alam, G. Aldering, D. M. Alexander, R. Alfarsy, L. Allen, C. Allende Prieto, O. Alves, A. Anand, U. Andrade *et al.* (DESI Collaboration), [arXiv:2503.14745](#).
- [27] A. G. Adame, J. Aguilar, S. Ahlen, S. Alam, D. M. Alexander, M. Alvarez, O. Alves, A. Anand, U. Andrade, E. Armengaud, S. Avila, A. Aviles, H. Awan, S. Bailey *et al.* (DESI Collaboration), *J. Cosmol. Astropart. Phys.* **04** (2025) 012.
- [28] A. G. Adame, J. Aguilar, S. Ahlen, S. Alam, D. M. Alexander, M. Alvarez, O. Alves, A. Anand, U. Andrade, E. Armengaud, S. Avila, A. Aviles, H. Awan, S. Bailey *et al.* (DESI Collaboration), *J. Cosmol. Astropart. Phys.* **01** (2025) 124.
- [29] A. G. Adame, J. Aguilar, S. Ahlen, S. Alam, D. M. Alexander, M. Alvarez, O. Alves, A. Anand, U. Andrade, E. Armengaud, S. Avila, A. Aviles, H. Awan, B. Bahr-Kalus *et al.* (DESI Collaboration), *J. Cosmol. Astropart. Phys.* **02** (2025) 021.
- [30] A. G. Adame, J. Aguilar, S. Ahlen, S. Alam, D. M. Alexander, M. Alvarez, O. Alves, A. Anand, U. Andrade, E. Armengaud, S. Avila, A. Aviles, H. Awan, S. Bailey *et al.* (DESI Collaboration), [arXiv:2411.12021](#).
- [31] A. G. Adame, J. Aguilar, S. Ahlen, S. Alam, D. M. Alexander, C. Allende Prieto, M. Alvarez, O. Alves, A. Anand, U. Andrade, E. Armengaud, S. Avila, A. Aviles, H. Awan *et al.* (DESI Collaboration), [arXiv:2411.12022](#).
- [32] M. Ishak, J. Pan, R. Calderon, K. Lodha, G. Valogiannis, A. Aviles, G. Niz, L. Yi, C. Zheng, C. Garcia-Quintero, A. de Mattia, L. Medina-Varela, J. L. Cervantes-Cota, U. Andrade, D. Huterer *et al.*, [arXiv:2411.12026](#).
- [33] M. Ishak *et al.*, [arXiv:2411.12026](#).
- [34] U. Andrade, J. Mena-Fernández, H. Awan, A. J. Ross, S. Brieden, J. Pan, A. de Mattia, J. Aguilar, S. Ahlen, O. Alves, D. Brooks, E. Buckley-Geer, E. Chaussidon, T. Claybaugh, S. Cole *et al.*, *J. Cosmol. Astropart. Phys.* **01** (2025) 128.
- [35] S. Brieden, H. Gil-Marín, L. Verde, and J. L. Bernal, *J. Cosmol. Astropart. Phys.* **09** (2020) 052.
- [36] E. Chaussidon, A. de Mattia, C. Yèche, J. Aguilar, S. Ahlen, D. Brooks, T. Claybaugh, S. Cole, A. de la Macorra, P. Doel, K. Fanning, E. Gaztañaga, S. G. A. Gontcho, C. Howlett, T. Kisner *et al.*, *J. Cosmol. Astropart. Phys.* **01** (2025) 135.
- [37] M. Abdul-Karim, J. Aguilar, S. Ahlen, S. Alam, L. Allen, C. Allende Prieto, O. Alves, A. Anand, U. Andrade, E. Armengaud, A. Aviles, S. Bailey, C. Baltay, P. Bansal *et al.* (DESI Collaboration), companion paper, *Phys. Rev. D* **112**, 083515 (2025).
- [38] A. Aghamousa, J. Aguilar, S. Ahlen, S. Alam, L. E. Allen, C. Allende Prieto, J. Annis, S. Bailey, C. Balland, O. Ballester, C. Baltay, L. Beaufore, C. Bebek, T. C. Beers *et al.* (DESI Collaboration), [arXiv:1611.00037](#).
- [39] B. Abareshi, J. Aguilar, S. Ahlen, S. Alam, D. M. Alexander, R. Alfarsy, L. Allen, C. Allende Prieto, O. Alves, J. Ameel, E. Armengaud, J. Asorey, A. Aviles, S. Bailey *et al.* (DESI Collaboration), *Astron. J.* **164**, 207 (2022).
- [40] A. D. Myers, J. Moustakas, S. Bailey, B. A. Weaver, A. P. Cooper, J. E. Forero-Romero, B. Abolfathi, D. M. Alexander, D. Brooks, E. Chaussidon, C.-H. Chuang, K. Dawson, A. Dey, B. Dey, G. Dhungana *et al.*, *Astron. J.* **165**, 50 (2023).
- [41] T. N. Miller, P. Doel, G. Gutierrez, R. Besuner, D. Brooks, G. Gallo, H. Heetderks, P. Jelinsky, S. M. Kent, M. Lampton, M. E. Levi, M. Liang, A. Meisner, M. J. Sholl, J. H. Silber *et al.*, *Astron. J.* **168**, 95 (2024).
- [42] J. H. Silber, P. Fagrelus, K. Fanning, M. Schubnell, J. N. Aguilar, S. Ahlen, J. Ameel, O. Ballester, C. Baltay, C. Bebek, D. Benton Beard, R. Besuner, L. Cardiel-Sas, R. Casas, F. J. Castander *et al.*, *Astron. J.* **165**, 9 (2023).
- [43] C. Poppett, L. Tyas, J. Aguilar, C. Bebek, D. Bramall, T. Claybaugh, J. Edelstein, P. Fagrelus, H. Heetderks, P. Jelinsky, S. Jelinsky, R. Lafever, A. Lambert, M. Lampton, M. E. Levi *et al.*, *Astron. J.* **168**, 245 (2024).
- [44] E. F. Schlafly, D. Kirkby, D. J. Schlegel, A. D. Myers, A. Raichoor, K. Dawson, J. Aguilar, C. Allende Prieto, S. Bailey, S. BenZvi, J. Bermejo-Climent, D. Brooks, A. de la Macorra, A. Dey, P. Doel *et al.*, *Astron. J.* **166**, 259 (2023).
- [45] J. Guy, S. Bailey, A. Kremin, S. Alam, D. M. Alexander, C. Allende Prieto, S. BenZvi, A. S. Bolton, D. Brooks, E. Chaussidon, A. P. Cooper, K. Dawson, A. de la Macorra, A. Dey, B. Dey *et al.*, *Astron. J.* **165**, 144 (2023).
- [46] Bailey *et al.* (to be published).
- [47] A. Anand *et al.* (DESI Collaboration), *Astron. J.* **168**, 124 (2024).
- [48] A. J. Ross, J. Aguilar, S. Ahlen, S. Alam, A. Anand, S. Bailey, D. Bianchi, S. Brieden, D. Brooks, E. Burtin, A. Carnero Rosell, E. Chaussidon, T. Claybaugh, S. Cole, K. Dawson *et al.*, *J. Cosmol. Astropart. Phys.* **01** (2025) 125.
- [49] A. G. Adame, J. Aguilar, S. Ahlen, S. Alam, D. M. Alexander, M. Alvarez, O. Alves, A. Anand, U. Andrade, E. Armengaud, S. Avila, A. Aviles, H. Awan, S. Bailey *et al.* (DESI Collaboration), [arXiv:2411.12020](#).
- [50] C. Hahn, M. J. Wilson, O. Ruiz-Macias, S. Cole, D. H. Weinberg, J. Moustakas, A. Kremin, J. L. Tinker, A. Smith, R. H. Wechsler, S. Ahlen, S. Alam, S. Bailey, D. Brooks, A. P. Cooper *et al.*, *Astron. J.* **165**, 253 (2023).
- [51] R. Zhou, B. Dey, J. A. Newman, D. J. Eisenstein, K. Dawson, S. Bailey, A. Berti, J. Guy, T.-W. Lan, H. Zou, J. Aguilar, S. Ahlen, S. Alam, D. Brooks, A. de la Macorra *et al.*, *Astron. J.* **165**, 58 (2023).

- [52] A. Raichoor, J. Moustakas, J. A. Newman, T. Karim, S. Ahlen, S. Alam, S. Bailey, D. Brooks, K. Dawson, A. de la Macorra, A. de Mattia, A. Dey, B. Dey, G. Dhungana, S. Eftekharzadeh *et al.*, *Astron. J.* **165**, 126 (2023).
- [53] E. Chaussidon, C. Yèche, N. Palanque-Delabrouille, D. M. Alexander, J. Yang, S. Ahlen, S. Bailey, D. Brooks, Z. Cai, S. Chabanier, T. M. Davis, K. Dawson, A. de laMacorra, A. Dey, B. Dey *et al.*, *Astrophys. J.* **944**, 107 (2023).
- [54] D. Valcin *et al.*, [arXiv:2508.05467](#).
- [55] A. J. Rosado-Marín, A. J. Ross, H. Seo, M. Rezaie, H. Kong, A. de Mattia, R. Zhou, S. Ahlen, D. Bianchi, D. Brooks, T. Claybaugh, A. de la Macorra, P. Doel, K. Fanning, S. Ferraro *et al.*, [arXiv:2411.12024](#).
- [56] H. Zou, X. Zhou, X. Fan, T. Zhang, Z. Zhou, J. Nie, X. Peng, I. McGreer, L. Jiang, A. Dey, D. Fan, B. He, Z. Jiang, D. Lang, M. Lesser *et al.*, *Publ. Astron. Soc. Pac.* **129**, 064101 (2017).
- [57] B. Flaugher, H. T. Diehl, K. Honscheid, T. M. C. Abbott, O. Alvarez, R. Angstadt, J. T. Annis, M. Antonik, O. Ballester, L. Beaufore, G. M. Bernstein, R. A. Bernstein, B. Bigelow, M. Bonati, D. Boprie *et al.*, *Astron. J.* **150**, 150 (2015).
- [58] T. Abbott, F. B. Abdalla, J. Aleksić, S. Allam, A. Amara, D. Bacon, E. Balbinot, M. Banerji, K. Bechtol, A. Benoit-Lévy, G. M. Bernstein, E. Bertin, J. Blazek, C. Bonnett *et al.* (Dark Energy Survey Collaboration), *Mon. Not. R. Astron. Soc.* **460**, 1270 (2016).
- [59] L. Breiman, *Mach. Learn.* **45**, 5 (2001).
- [60] K. Bundy, A. Leauthaud, S. Saito, A. Bolton, Y.-T. Lin, C. Maraston, R. C. Nichol, D. P. Schneider, D. Thomas, and D. A. Wake, *Astrophys. J. Suppl. Ser.* **221**, 15 (2015).
- [61] H. A. Feldman, N. Kaiser, and J. A. Peacock, *Astrophys. J.* **426**, 23 (1994).
- [62] S. D. Landy and A. S. Szalay, *Astrophys. J.* **412**, 64 (1993).
- [63] N. Padmanabhan, X. Xu, D. J. Eisenstein, R. Scalzo, A. J. Cuesta, K. T. Mehta, and E. Kazin, *Mon. Not. R. Astron. Soc.* **427**, 2132 (2012).
- [64] M. Sinha and L. H. Garrison, *Mon. Not. R. Astron. Soc.* **491**, 3022 (2020).
- [65] M. White, *Mon. Not. R. Astron. Soc.* **450**, 3822 (2015).
- [66] K. Yamamoto, M. Nakamichi, A. Kamino, B. A. Bassett, and H. Nishioka, *Publ. Astron. Soc. Jpn.* **58**, 93 (2006).
- [67] D. Bianchi, H. Gil-Marín, R. Ruggeri, and W. J. Percival, *Mon. Not. R. Astron. Soc.* **453**, L11 (2015).
- [68] N. Hand, Y. Feng, F. Beutler, Y. Li, C. Modi, U. Seljak, and Z. Slepian, *Astron. J.* **156**, 160 (2018).
- [69] N. Dalal, O. Dore, D. Huterer, and A. Shirokov, *Phys. Rev. D* **77**, 123514 (2008).
- [70] N. Aghanim, Y. Akrami, M. Ashdown, J. Aumont, C. Baccigalupi, M. Ballardini, A. J. Banday, R. B. Barreiro, N. Bartolo, S. Basak, R. Battye, K. Benabed, J. P. Bernard, M. Bersanelli *et al.* (Planck Collaboration), *Astron. Astrophys.* **641**, A6 (2020).
- [71] M. Chevallier and D. Polarski, *Int. J. Mod. Phys. D* **10**, 213 (2001).
- [72] E. V. Linder, *Phys. Rev. Lett.* **90**, 091301 (2003).
- [73] J. Lasker, A. Carnero Rosell, A. D. Myers, A. J. Ross, D. Bianchi, M. M. S. Hanif, R. Kehoe, A. de Mattia, L. Napolitano, W. J. Percival, R. Staten, J. Aguilar, S. Ahlen, L. Bigwood, D. Brooks *et al.*, *J. Cosmol. Astropart. Phys.* **01** (2025) 127.
- [74] C. Zhao *et al.* (to be published).
- [75] N. A. Maksimova, L. H. Garrison, D. J. Eisenstein, B. Hadzhiyska, S. Bose, and T. P. Satterthwaite, *Mon. Not. R. Astron. Soc.* **508**, 4017 (2021).
- [76] S. Yuan, H. Zhang, A. J. Ross, J. Donald-McCann, B. Hadzhiyska, R. H. Wechsler, Z. Zheng, S. Alam, V. Gonzalez-Perez, J. N. Aguilar, S. Ahlen, D. Bianchi, D. Brooks, A. de la Macorra, K. Fanning *et al.*, [arXiv:2306.06314](#).
- [77] A. Smith, C. Grove, S. Cole, P. Norberg, P. Zarrouk, S. Yuan, J. Aguilar, S. Ahlen, D. Brooks, T. Claybaugh, A. de la Macorra, P. Doel, J. E. Forero-Romero, E. Gaztañaga, S. G. A. Gontcho *et al.*, *Mon. Not. R. Astron. Soc.* **532**, 903 (2024).
- [78] A. Rocher, V. Ruhlmann-Kleider, E. Burtin, S. Yuan, A. de Mattia, A. J. Ross, J. Aguilar, S. Ahlen, S. Alam, D. Bianchi, D. Brooks, S. Cole, K. Dawson, A. de la Macorra, P. Doel *et al.*, *J. Cosmol. Astropart. Phys.* **10** (2023) 016.
- [79] S. Yuan, H. Zhang, A. J. Ross, J. Donald-McCann, B. Hadzhiyska, R. H. Wechsler, Z. Zheng, S. Alam, V. Gonzalez-Perez, J. N. Aguilar, S. Ahlen, D. Bianchi, D. Brooks, A. de la Macorra, K. Fanning *et al.*, *Mon. Not. R. Astron. Soc.* **530**, 947 (2024).
- [80] A. Smith, C. Grove, S. Cole, P. Norberg, P. Zarrouk, S. Yuan, J. Aguilar, S. Ahlen, D. Brooks, T. Claybaugh, A. de la Macorra, P. Doel, J. E. Forero-Romero, E. Gaztañaga, S. G. A. Gontcho *et al.*, *Mon. Not. R. Astron. Soc.* **532**, 903 (2024).
- [81] C. Alcock and B. Paczynski, *Nature (London)* **281**, 358 (1979).
- [82] D. J. Eisenstein, H.-J. Seo, E. Sirko, and D. N. Spergel, *Astrophys. J.* **664**, 675 (2007).
- [83] S. F. Chen, C. Howlett, M. White, P. McDonald, A. J. Ross, H. J. Seo, N. Padmanabhan, J. Aguilar, S. Ahlen, S. Alam, O. Alves, U. Andrade, R. Blum, D. Brooks, X. Chen *et al.*, *Mon. Not. R. Astron. Soc.* **534**, 544 (2024).
- [84] X. Chen, Z. Ding, E. Paillas, S. Nadathur, H. Seo, S. Chen, N. Padmanabhan, M. White, A. de Mattia, P. McDonald, A. J. Ross, A. Variu, A. Carnero Rosell, B. Hadzhiyska, M. M. S. Hanif *et al.*, [arXiv:2411.19738](#).
- [85] A. Burden, W. J. Percival, and C. Howlett, *Mon. Not. R. Astron. Soc.* **453**, 456 (2015).
- [86] E. Paillas, Z. Ding, X. Chen, H. Seo, N. Padmanabhan, A. de Mattia, A. J. Ross, S. Nadathur, C. Howlett, J. Aguilar, S. Ahlen, O. Alves, U. Andrade, D. Brooks, E. Buckley-Geer *et al.*, *J. Cosmol. Astropart. Phys.* **01** (2025) 142.
- [87] C. Garcia-Quintero, J. Mena-Fernández, A. Rocher, S. Yuan, B. Hadzhiyska, O. Alves, M. Rashkovetskyi, H. Seo, N. Padmanabhan, S. Nadathur, C. Howlett, M. Ishak, L. Medina-Varela, P. McDonald, A. J. Ross *et al.*, *J. Cosmol. Astropart. Phys.* **01** (2025) 132.
- [88] J. Mena-Fernández, C. Garcia-Quintero, S. Yuan, B. Hadzhiyska, O. Alves, M. Rashkovetskyi, H. Seo, N. Padmanabhan, S. Nadathur, C. Howlett, S. Alam, A. Rocher, A. J. Ross, E. Sanchez, M. Ishak *et al.*, *J. Cosmol. Astropart. Phys.* **01** (2025) 133.
- [89] A. Pérez-Fernández, L. Medina-Varela, R. Ruggeri, M. Vargas-Magaña, H. Seo, N. Padmanabhan, M. Ishak,

- J. Aguilar, S. Ahlen, S. Alam, O. Alves, U. Andrade, S. Brieden, D. Brooks, A. Carnero Rosell *et al.*, *J. Cosmol. Astropart. Phys.* **01** (2025) 144.
- [90] S. Brieden, H. Gil-Marín, and L. Verde, *J. Cosmol. Astropart. Phys.* **08** (2022) 024.
- [91] M. Pinon, A. de Mattia, P. McDonald, E. Burtin, V. Ruhlmann-Kleider, M. White, D. Bianchi, A. J. Ross, J. Aguilar, S. Ahlen, D. Brooks, R. N. Cahn, E. Chaussidon, T. Claybaugh, S. Cole *et al.*, *J. Cosmol. Astropart. Phys.* **01** (2025) 131.
- [92] R. O’Connell, D. Eisenstein, M. Vargas, S. Ho, and N. Padmanabhan, *Mon. Not. R. Astron. Soc.* **462**, 2681 (2016).
- [93] R. O’Connell and D. J. Eisenstein, *Mon. Not. R. Astron. Soc.* **487**, 2701 (2019).
- [94] O. H. E. Philcox, D. J. Eisenstein, R. O’Connell, and A. Wiegand, *Mon. Not. R. Astron. Soc.* **491**, 3290 (2020).
- [95] O. H. E. Philcox and D. J. Eisenstein, *Mon. Not. R. Astron. Soc.* **490**, 5931 (2019).
- [96] M. Rashkovetskyi, D. J. Eisenstein, J. N. Aguilar, D. Brooks, T. Claybaugh, S. Cole, K. Dawson, A. de la Macorra, P. Doel, K. Fanning, A. Font-Ribera, J. E. Forero-Romero, S. Gontcho A Gontcho, C. Hahn, K. Honscheid *et al.*, *Mon. Not. R. Astron. Soc.* **524**, 3894 (2023).
- [97] M. Rashkovetskyi, D. Forero-Sánchez, A. de Mattia, D. J. Eisenstein, N. Padmanabhan, H. Seo, A. J. Ross, J. Aguilar, S. Ahlen, O. Alves, U. Andrade, D. Brooks, E. Burtin, X. Chen, T. Claybaugh *et al.*, *J. Cosmol. Astropart. Phys.* **01** (2025) 145.
- [98] D. Wadekar and R. Scoccimarro, *Phys. Rev. D* **102**, 123517 (2020).
- [99] O. Alves *et al.* (to be published).
- [100] R. Zhao, K. Koyama, Y. Wang, and G.-B. Zhao, *Res. Astron. Astrophys.* **24**, 125015 (2024).
- [101] D. Foreman-Mackey, D. W. Hogg, D. Lang, and J. Goodman, *Publ. Astron. Soc. Pac.* **125**, 306 (2013).
- [102] A. Gelman and D. B. Rubin, *Stat. Sci.* **7**, 457 (1992).
- [103] H. Dembinski, P. Ongmongkolkul *et al.*, [10.5281/zenodo.3949207](https://arxiv.org/abs/10.5281/zenodo.3949207) (2020).
- [104] Sanders *et al.* (to be published).
- [105] A. Krolewski, J. Yu, A. J. Ross, S. Penmetza, W. J. Percival, R. Zhou, M. J. Wilson, J. Hou, J. Aguilar, S. Ahlen, D. Brooks, E. Chaussidon, T. Claybaugh, A. de la Macorra, B. Dey *et al.*, *J. Cosmol. Astropart. Phys.* **01** (2025) 147.
- [106] J. Yu, A. J. Ross, A. Rocher, O. Alves, A. de Mattia, D. Forero-Sánchez, J. Kneib, A. Krolewski, T. W. Lan, M. Rashkovetskyi, J. Aguilar, S. Ahlen, S. Bailey, D. Brooks, E. Chaussidon *et al.*, *J. Cosmol. Astropart. Phys.* **01** (2025) 126.
- [107] S. E. Koposov *et al.*, *Mon. Not. R. Astron. Soc.* **533**, 1012 (2024).
- [108] N. Aghanim, Y. Akrami, F. Arroja, M. Ashdown, J. Aumont, C. Baccigalupi, M. Ballardini, A. J. Banday, R. B. Barreiro, N. Bartolo, S. Basak, R. Battye, K. Benabed, J. P. Bernard *et al.* (Planck Collaboration), *Astron. Astrophys.* **641**, A1 (2020).
- [109] J. E. Bautista, R. Paviot, M. Vargas Magaña, S. de la Torre, S. Fromenteau, H. Gil-Marín, A. J. Ross, E. Burtin, K. S. Dawson, J. Hou, J.-P. Kneib, A. de Mattia, W. J. Percival, G. Rossi, R. Tojeiro *et al.*, *Mon. Not. R. Astron. Soc.* **500**, 736 (2021).
- [110] A. J. Ross, F. Beutler, C.-H. Chuang, M. Pellejero-Ibanez, H.-J. Seo, M. Vargas-Magaña, A. J. Cuesta, W. J. Percival, A. Burden, A. G. Sánchez, J. N. Grieb, B. Reid, J. R. Brownstein, K. S. Dawson, D. J. Eisenstein *et al.*, *Mon. Not. R. Astron. Soc.* **464**, 1168 (2017).
- [111] C. Saulder, Y.-S. Song, M. Oh, Y. Zheng, A. J. Ross, R. Zhou, J. A. Newman, C.-H. Chuang, J. N. Aguilar, S. Ahlen, R. Blum, D. Brooks, T. Claybaugh, A. de la Macorra, B. Dey *et al.*, *Astron. Astrophys.* **695**, A54 (2025).
- [112] A. G. Adame, J. Aguilar, S. Ahlen, S. Alam, G. Aldering, D. M. Alexander, R. Alfarsy, C. Allende Prieto, M. Alvarez, O. Alves, A. Anand, F. Andrade-Oliveira, E. Armengaud, J. Asorey *et al.* (DESI Collaboration), *Astron. J.* **167**, 62 (2024).
- [113] J. Pan, D. Huterer, F. Andrade-Oliveira, and C. Avestruz, *J. Cosmol. Astropart. Phys.* **06** (2024) 051.
- [114] J. L. Bernal, T. L. Smith, K. K. Boddy, and M. Kamionkowski, *Phys. Rev. D* **102**, 123515 (2020).
- [115] K. Thepsuriya and A. Lewis, *J. Cosmol. Astropart. Phys.* **01** (2015) 034.
- [116] M. Abdul-Karim, J. Aguilar, S. Ahlen, C. Allende Prieto, O. Alves, A. Anand, U. Andrade, E. Armengaud, A. Aviles, S. Bailey, A. Bault, S. BenZvi, D. Bianchi, C. Blake *et al.* (DESI Collaboration), [arXiv:2503.14739](https://arxiv.org/abs/2503.14739).
- [117] <https://www.desi.lbl.gov/collaborating-institutions>.
- [118] <https://data.desi.lbl.gov/doc/releases/>.
- [119] H.-J. Seo and D. J. Eisenstein, *Astrophys. J.* **665**, 14 (2007).